

PAPER • OPEN ACCESS

Studying dark matter with MadDM: lines and loops

To cite this article: D. Massaro *et al* 2021 *J. Phys.: Conf. Ser.* **2156** 012073

View the [article online](#) for updates and enhancements.

You may also like

- [In situ Optical Emission Spectrometry during the Initiation and Growth of Porous Anodic Alumina](#)
Dimitri Mercier, Quentin Van Overmeere, Ronny Santoro *et al.*
- [Temperature dependence of penetration and coherence lengths in lead nanowires](#)
G Stenuit, S Michotte, J Govaerts *et al.*
- [Topological analysis of traffic pace via persistent homology](#)
Daniel R Carmody and Richard B Sowers



The Electrochemical Society
Advancing solid state & electrochemical science & technology

242nd ECS Meeting

Oct 9 – 13, 2022 • Atlanta, GA, US

Abstract submission deadline: **April 8, 2022**

Connect. Engage. Champion. Empower. Accelerate.

MOVE SCIENCE FORWARD



Submit your abstract



Studying dark matter with MADDM: lines and loops

D. Massaro^{1,2,3}, C. Arina⁴, J. Heisig^{5,3}, F. Maltoni^{3,1,2} and O. Mattelaer³

¹ Dipartimento di Fisica e Astronomia, Alma Mater Studiorum - Università di Bologna, via Irnerio 46, 40126 Bologna, Italy

² INFN, Sezione di Bologna, viale Berti Pichat 6/2, 40127 Bologna, Italy

³ Centre for Cosmology, Particle Physics and Phenomenology (CP3), Université catholique de Louvain, Chemin du Cyclotron 2, B-1348 Louvain-la-Neuve, Belgium

⁴ Research Institute for Mathematics and Physics (IRMP), Université catholique de Louvain, Chemin du Cyclotron 2, B-1348 Louvain-la-Neuve, Belgium

⁵ Institute for Theoretical Particle Physics and Cosmology, RWTH Aachen University, Sommerfeldstr. 16, D-52056 Aachen, Germany

E-mail: daniele.massaro5@unibo.it

Abstract. Automated tools for the computation of amplitudes and cross sections have become the backbone of phenomenological studies beyond the standard model. We present the latest developments in MADDM, a calculator of dark-matter observables based on MADGRAPH5_AMC@NLO. The new version enables the fully automated computation of loop-induced annihilation processes, relevant for indirect detection of dark matter. Of particular interest is the direct annihilation into photons, $\gamma\gamma$, γX . These processes lead to monochromatic gamma-ray lines that are smoking-gun signatures for dark-matter annihilation in our Galaxy. MADDM computes the predictions for the expected photon fluxes near Earth and derives constraints from the gamma-ray line searches by Fermi-LAT and HESS. As an application, we present the implications for the parameter space of the Inert Doublet Model.

1. Introduction

The existence of dark matter has been corroborated by numerous observations over a wide range of physical scales in the Universe. Indirect detection is one of the main search strategies to shed light on its nature. It explores traces of dark-matter annihilation (or decay) in cosmic messengers. In this context, direct annihilation into neutral particles can give rise to prominent spectral features, such as monochromatic lines. These are smoking-gun signals for dark-matter annihilation, as the pronounced peak is hardly mimicked by any astrophysical background. In this report, we consider direct annihilation into photons, i.e. $\gamma\gamma$ and γX , where $X = Z, h$ or any new particle even under the dark symmetry. These are loop-induced processes given that dark matter is electrically neutral. The experimental sensitivity to such sharp photon energy spectra is high: current gamma-ray telescopes, like Fermi-LAT [1] and HESS [2], are able to set strong constraints on the dark-matter parameter space.

Here, we present MADDM v3.2 [3], that features the automatized computation of loop-induced processes for indirect detection within any dark-matter model for which a UFO [4] model at next-to-leading order (NLO) can be generated. At the time of writing, it is the only numerical tool with such capability. It further includes the automated computation of the



photon fluxes for a variety of dark-matter density profiles and the derivation of exclusion limits by confronting the predictions with the Fermi-LAT satellite and the HESS telescope upper limits. This new module builds upon the indirect detection module, released with MADDM v3.0 [5].

This report is organised as follows. In § 2 we describe the relevant astrophysical aspects of gamma-ray line searches, in § 3 we outline the main functionalities of the new MADDM version, while § 4 provides a phenomenological study of gamma-ray line signatures within the Inert Doublet Model (IDM). We conclude in § 5.

2. Gamma-ray line phenomenology

2.1. Gamma-ray flux and photon energy

The differential flux of gamma rays from annihilation of dark matter in a given region of interest (ROI) is given by¹

$$\frac{d\Phi}{dE} = \frac{1}{8\pi m_{\text{DM}}^2} \sum_i \langle \sigma v \rangle_i \frac{dN_\gamma^i}{dE} \underbrace{\int_{\text{ROI}} d\Omega \int_{\text{l.o.s.}} \rho^2(\mathbf{r}) dl}_{J\text{-factor}}, \quad (1)$$

where $\langle \sigma v \rangle_i$ is the velocity averaged cross-section of dark-matter particles with a mass m_{DM} into final states labelled by i and dN_γ^i/dE is the respective differential gamma-ray energy spectrum per annihilation. The second part of the equation defines the J -factor, where $\rho(\mathbf{r})$ denotes the dark matter density distribution. The second integral is performed over the line of sight (l.o.s.) l . In MADDM, the following dark-matter density profiles are implemented: the Navarro-Frenk-White (NFW) [6], Einasto [7], Burkert [8] and isothermal profile [9].

Dark matter is assumed to be non-relativistic in galactic halos: typical values of the dark-matter velocity within the Milky Way [10], and for similar galaxies, are of the order of $v \simeq 10^{-3}c$. Accordingly, for dark-matter annihilation into $\gamma\gamma$ or γX , the differential gamma-ray energy spectrum is:

$$\frac{dN_\gamma}{dE} = \delta(E - E_\gamma) \times \begin{cases} 2 & \text{for } \gamma\gamma \\ 1 & \text{for } \gamma X \end{cases}, \quad (2)$$

where X denotes a neutral standard model or beyond standard model (BSM) particle with mass m_X and

$$E_\gamma = m_{\text{DM}} \left(1 - \frac{m_X^2}{4m_{\text{DM}}^2} \right). \quad (3)$$

This characteristic shape allows for peak searches within the experimental data.

2.2. Experimental constraints

Searches for gamma-ray line signals from dark-matter annihilation from the galactic centre have been performed, for instance, by the Fermi-LAT satellite [1] and the HESS telescope [2]. Fermi-LAT data collected over 5.8 years of observation, provide upper limits in the annihilation cross-section into di-photons constraining dark-matter masses in the range 200 MeV to 500 GeV. The data collected by the HESS telescope are based on 254 h of live-time observation and are able to constrain rather heavy dark-matter masses, from 300 GeV up to a maximum of 70 TeV.

In the case of multiple final states, the application of the experimental constraints requires particular attention. The photon energy spectrum is a superposition of lines (which in the following will be called *peaks*) coming from the various final states $\gamma\gamma$, γZ , γh , according to eq. (3). In the experimental measurement, the predicted gamma-line, a Dirac delta function, is

¹ For non self-conjugate dark-matter candidates eq. (1) has to be multiplied by an additional factor of 1/2.

smearred into a Gaussian distribution according to the energy resolution, so that peaks sufficiently close in energy could be indistinguishable for an experiment. In this case, they should be combined and the final peak's flux would be the sum of the fluxes of each peak. Furthermore, (combined) spectral lines lacking a sufficient separation from each other question the applicability of the experimental limit-setting procedure, because the experimental analysis has been carried out in the hypothesis of a single photon line.

3. Loop-induced processes in MadDM

MADLOOP [11] has been used extensively to compute loop-induced processes in the framework of collider searches for dark matter within MG5_AMC [12], and can now be used within MADDM v3.2 to compute loop-induced annihilation processes in galactic halos. To be able to perform loop computations within MG5_AMC and MADDM it is necessary to import NLO UFO model files. This can be achieved by using FEYNRULES [13], FEYNARTS [14] and NLOCT [15].

Then, for a given NLO UFO dark-matter model, MADDM automatically generates all contributing diagrams for annihilation into γX with the command:

```
MadDM> generate indirect_spectral_features
```

where X includes the Z and h particles of the standard model as well as *all* additional BSM particles, that are lighter than twice the dark-matter mass and transform even under the dark symmetry that stabilizes the dark matter. Individual channels can be generated by explicitly specifying the final state, e.g.

```
MadDM> generate indirect_spectral_features a z
```

The analysis pipeline for these final states starts with the computation of the annihilation cross-section and of the J -factor. Subsequently, it performs a combination (if any) of all peaks to obtain the full gamma-line energy spectrum, finally yielding the prediction of the γ flux, and compares it against the current experimental constraints, see § 2.2.

Loop-induced annihilation into final states other than the photon ones may contribute to the continuum flux of cosmic messengers. A relevant example is the dark-matter annihilation into a pair of gluons, gg , that subsequently shower and hadronize. In the absence of a tree-level diagram for this channel, MADDM automatically switches to the loop-induced mode. Hence, this annihilation channel is considered by executing the command:

```
MadDM> generate indirect_detection g g
```

It can also be computed together with tree-level diagrams:

```
MadDM> generate indirect_detection
MadDM> add indirect_detection g g
```

Here, the first line leads to the computation of all $2 \rightarrow 2$ tree-level annihilation processes. After the computation of the annihilation cross-section and generation of events, MADDM proceeds with the indirect detection analysis pipeline as introduced in [5].

4. Application to the IDM

To demonstrate the physics impact, in the following we apply the new feature of MADDM to the IDM, that is built on top of the standard model by adding a new (*inert*) Higgs doublet, Φ , odd under an exact Z_2 symmetry. The scalar potential reads

$$V = \mu_1^2 |H|^2 + \mu_2^2 |\Phi|^2 + \lambda_1 |H|^4 + \lambda_2 |\Phi|^4 + \lambda_3 |H|^2 |\Phi|^2 + \lambda_4 |H^\dagger \Phi|^2 + \frac{\lambda_5}{2} [(H^\dagger \Phi)^2 + \text{h.c.}]. \quad (4)$$

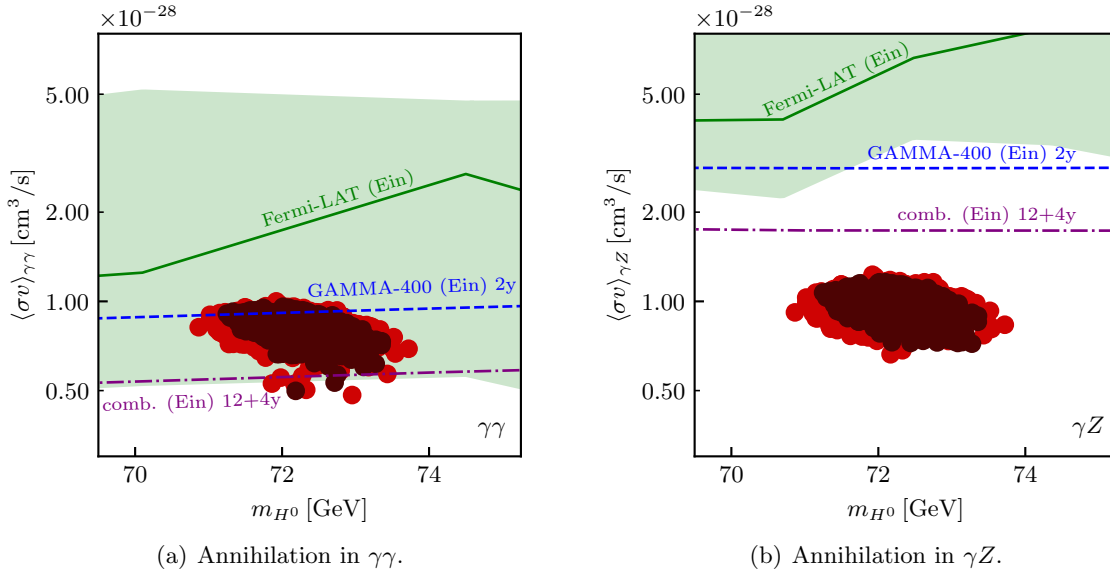


Figure 1. Cross section for the loop-induced dark-matter annihilation for the allowed IDM parameter points explaining all of dark matter ($\Omega h^2 = 0.12$) in the range $70 \sim 75$ GeV. Parameter points in the 1 and 2σ regions are drawn in dark brown and red full circles, respectively. The green solid curve denotes the Fermi-LAT limit for the Einasto profile. The lower and upper boundary of the green shaded band shows the respective limits for NFW contracted and isothermal profile, respectively. The blue dotted and purple dot-dashed curves show projections for GAMMA-400 (2 years) and the combination of Fermi-LAT (12 years) and GAMMA-400 (4 years), respectively, both assuming the Einasto profile.

After electroweak symmetry breaking, with $\Phi = (H^\pm, (H^0 + iA^0)/\sqrt{2})^T$, we obtain a total of five physical scalar states: h, H^0, A^0, H^\pm . We are left with five free parameters: $m_{H^0}, m_{A^0}, m_{H^\pm}, \lambda_L, \lambda_2$, where $\lambda_{L,S} = (\lambda_3 + \lambda_4 \pm \lambda_5)/2$. Here, we assume H^0 to be the dark-matter candidate and focus on the phenomenologically interesting region around $m_{H^0} \simeq 72$ GeV. In this region, the measured relic density can be explained by annihilation into WW^*, ZZ^* via the gauge kinetic interaction alone. (Here, V^* denotes an off-shell vector boson.) The region is currently unchallenged by current limits from the LHC and direct detection [16].

We generate the NLO UFO model with FEYNRULES, FEYNARTS and NLOCT and consider the processes $H^0 H^0 \rightarrow \gamma\gamma$ and $H^0 H^0 \rightarrow \gamma Z$ involving 140 and 172 diagrams, respectively. Note that $H^0 H^0 \rightarrow \gamma h$ is forbidden due to charge-conjugation invariance. We have validated our numerical setup using existing results for $\langle\sigma v\rangle_{\gamma\gamma}$ in the literature [17, 18] and found agreement within the numerical precision. We employ the parameter scan performed in [16] and consider points within the 2σ -region taking into account constraints from the relic density, electroweak precision observables, new physics searches at LEP-II, indirect detection searches for continuous gamma-ray spectra and theoretical requirements of unitarity, perturbativity and vacuum stability. The resulting cross sections for $\langle\sigma v\rangle_{\gamma\gamma}$ and $\langle\sigma v\rangle_{\gamma Z}$ are shown in fig. 1 together with the corresponding upper limits from Fermi-LAT [1] and future projections for GAMMA-400 (2 years) as well as a combination of both observations (with 12 and 4 years observational time, respectively) [19]. All lines assume the Einasto dark-matter density profile while the shaded band around the Fermi-LAT limit illustrates the uncertainty due to the choice of the profile. The upper and lower boundaries of the shaded band corresponds to the upper limit assuming the isothermal and NFWc profiles, respectively. In the $\gamma\gamma$ channel, the current upper limits only constrain the considered region for the case of NFW (contracted) profile. For the Einasto

profile, future observations are expected to provide sensitivity. In the γZ channel, the sensitivity of the combination of Fermi-LAT and GAMMA-400 is still too low by about a factor of two to reach the expected signal originating.

5. Conclusion

Gamma-ray lines are smoking-gun signatures of direct dark-matter annihilation into photons in galactic halos. For electrically neutral dark matter, annihilation into $\gamma\gamma$ and γX is loop-induced. In this report, we presented MADDM v3.2, that enables the automated computation of loop-induced cross sections for arbitrary dark-matter models implemented in the NLO UFO format via an interface to MADLOOP. Furthermore, MADDM computes the resulting integrated photon fluxes and applies experimental constraints from Fermi-LAT and HESS. It also allows the user to perform the computation of the J -factor for a variety of dark-matter density profiles and ROIs.

We demonstrated the capabilities of the program by applying it to the IDM, concentrating on the phenomenologically interesting region of the parameter space $71 \text{ GeV} \lesssim m_{\text{DM}} \lesssim 74 \text{ GeV}$. We found that current limits from gamma-line observation constrain the scenario for the very cuspy NFW (contracted) profile while future observations by GAMMA-400 and Fermi-LAT are expected to probe it for a slightly broader range of dark-matter density profiles. An intriguing possibility pointing to the IDM model would be the observation of two peaks, from $\gamma\gamma$ and γZ , which are well separated for the considered dark-matter masses. However, the γZ signal still appears out-of-reach according to the above-mentioned projections.

References

- [1] Ackermann M *et al.* (Fermi-LAT) 2015 *Phys. Rev. D* **91** 122002 (*Preprint* 1506.00013)
- [2] Abdallah H *et al.* (HESS) 2018 *Phys. Rev. Lett.* **120** 201101 (*Preprint* 1805.05741)
- [3] Arina C, Heisig J, Maltoni F, Massaro D and Mattelaer O 2021 *Preprint* 2107.04598
- [4] Degrande C, Duhr C, Fuks B, Grellscheid D, Mattelaer O and Reiter T 2012 *Comput. Phys. Commun.* **183** 1201–1214 (*Preprint* 1108.2040)
- [5] Ambrogio F, Arina C, Backovic M, Heisig J, Maltoni F, Mantani L, Mattelaer O and Mohlabeng G 2019 *Phys. Dark Univ.* **24** 100249 (*Preprint* 1804.00044)
- [6] Navarro J F, Frenk C S and White S D M 1996 *Astrophys. J.* **462** 563–575 (*Preprint* astro-ph/9508025)
- [7] Einasto J 1965 *Trudy Astrofizicheskogo Instituta Alma-Ata* **5** 87–100
- [8] Burkert A 1995 *Astrophys. J. Lett.* **447** L25 (*Preprint* astro-ph/9504041)
- [9] Binney J and Tremaine S 2008 *Galactic Dynamics: Second Edition* (Princeton University Press)
- [10] Eilers A C, Hogg D W, Rix H W and Ness M K 2019 *The Astrophysical Journal* **871** 120 (*Preprint* 1810.09466)
- [11] Hirschi V, Frederix R, Frixione S, Garzelli M V, Maltoni F and Pittau R 2011 *JHEP* **05** 044 (*Preprint* 1103.0621)
- [12] Hirschi V and Mattelaer O 2015 *JHEP* **10** 146 (*Preprint* 1507.00020)
- [13] Alloul A, Christensen N D, Degrande C, Duhr C and Fuks B 2014 *Comput. Phys. Commun.* **185** 2250–2300 (*Preprint* 1310.1921)
- [14] Hahn T 2001 *Comput. Phys. Commun.* **140** 418–431 (*Preprint* hep-ph/0012260)
- [15] Degrande C 2015 *Comput. Phys. Commun.* **197** 239–262 (*Preprint* 1406.3030)
- [16] Eiteneuer B, Goudelis A and Heisig J 2017 *Eur. Phys. J. C* **77** 624 (*Preprint* 1705.01458)
- [17] Gustafsson M, Lundstrom E, Bergstrom L and Edsjo J 2007 *Phys. Rev. Lett.* **99** 041301 (*Preprint* astro-ph/0703512)
- [18] Garcia-Cely C and Rivera A 2017 *JCAP* **03** 054 (*Preprint* 1611.08029)
- [19] Egorov A E, Topchiev N P, Galper A M, Dalkarov O D, Leonov A A, Suchkov S I and Yurkin Y T 2020 *JCAP* **11** 049 (*Preprint* 2005.09032)