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Probing AGN inner structure with X-ray obscured type 1 AGN

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ABSTRACT

Using the X-ray-selected active galactic nuclei (AGN) from the *XMM-XXL* north survey and the Sloan Digital Sky Survey-Baryon Oscillation Spectroscopic Survey spectroscopic follow-up of them, we compare the properties of X-ray unobscured and obscured broad-line AGN (BLAGN1 and BLAGN2; N_{H} below and above $10^{21.5} \text{ cm}^{-2}$), including their X-ray luminosity L_{X} , black hole mass, Eddington ratio λ_{Edd} , optical continuum, and line features. We find that BLAGN2 have systematically larger broad line widths and hence apparently higher (lower) M_{BH} (λ_{Edd}) than BLAGN1. We also find that the X-ray obscuration in BLAGN tends to coincide with optical dust extinction, which is optically thinner than that in narrow-line AGN (NLAGN) and likely partial covering to the broad-line region. All the results can be explained in the framework of a multicomponent, clumpy torus model by interpreting BLAGN2 as an intermediate type between BLAGN1 and NLAGN in terms of an intermediate inclination angle.

Key words: surveys – galaxies: active – galaxies: nuclei – quasars: emission lines – X-rays: galaxies.

1 INTRODUCTION

Within the basic scheme of active galactic nuclei (AGN) unification model, both the differences between X-ray unobscured (X-ray type 1) and obscured (X-ray type 2) AGN and between broad-line (optical type 1) and narrow-line (optical type 2) AGN are determined by inclination angles with respect to an obscuring dusty ‘torus’. This axisymmetric ‘torus’ plays an essential role in the unification model (Antonucci 1993). However, even recent Atacama Large Millimeter/submillimeter Array (ALMA) high-resolution observations could only resolve a rotating circumnuclear disc for the nearby Seyfert galaxy NGC 1068 (García-Burillo et al. 2016; Imanishi et al. 2018). The detailed structure of the ‘torus’ is unclear, not to mention the physical mechanism that regulates it. Especially, for

a portion of AGN, the optical and X-ray classifications of type 1 and type 2 disagree with each other, complicating the understanding of the ‘torus’ (Brusa et al. 2003; Perola et al. 2004; Merloni et al. 2014; Davies et al. 2015).

A large amount of work has been devoted to the study of the correlation between X-ray obscuration and luminosity. Generally, relatively higher column densities (or larger obscured fractions) are found at lower luminosities (e.g. Lawrence & Elvis 1982; Treister & Urry 2006; Hasinger 2008; Brightman & Nandra 2011; Burlon et al. 2011; Lusso et al. 2013; Brightman et al. 2014). However, until reliable black hole masses (M_{BH}) for a sample of AGN are measured accurately, one could not clearly reveal the correlation between the obscuration and the M_{BH} -normalized accretion rate (i.e. the Eddington ratio, λ_{Edd}), which is considered as the main physical driver of the principal component of AGN properties (e.g. Boroson & Green 1992; Sulentic et al. 2000). Using the *Swift*-Burst Alert Telescope

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(BAT) selected local AGN sample, Ricci et al. (2017) found that the AGN obscured fraction is mainly determined by the λ_{Edd} rather than the luminosity, and concluded that the main physical driver of the torus diversity is λ_{Edd} , which regulates the torus covering factor by means of radiation pressure. To test the role of λ_{Edd} in regulating the obscuration of AGN, the M_{BH} of X-ray obscured and unobscured AGN must be measured consistently to avoid possible biases. Except for the tens of AGN whose M_{BH} could be measured by reverberation mapping or dynamical methods (e.g. Peterson et al. 2004), generally the M_{BH} of the X-ray unobscured AGN are measured on the basis of the broad line widths and the continuum luminosity (single-epoch virial method); for X-ray obscured AGN, the M_{BH} are inferred on the basis of the empirical relation between the M_{BH} and stellar velocity dispersion (e.g. Ho et al. 2012; Bisogni et al. 2017a; Koss et al. 2017). The X-ray obscuration presented in a small fraction of broad-line AGN (BLAGN) provides a great tool for this test, since the M_{BH} of the X-ray unobscured and obscured BLAGN (BLAGN1 and BLAGN2) can be measured consistently using the same method. Even then, the single-epoch M_{BH} must be used with caution, considering that the virial f factor can be inclination dependent (Wills & Browne 1986; Risaliti, Salvati & Marconi 2011; Pancoast et al. 2014; Shen & Ho 2014; Bisogni, Marconi & Risaliti 2017b; Mejía-Restrepo et al. 2018).

In BLAGN, whose broad-line region (BLR) is visible, it is unclear what causes the X-ray obscuration. The X-ray obscuring material might be dust-free and therefore transparent to optical emission from the accretion disc and BLR (Merloni et al. 2014; Davies et al. 2015; Liu et al. 2016); or it might be a dusty cloud blocking only the central engine (accretion disc and corona) but not the BLR because of geometric reasons, e.g. small obscuring cloud moving across the line of sight of the X-ray emitting corona (e.g. Risaliti, Elvis & Nicastro 2002; Maiolino et al. 2010). Study of multiband emission and obscuration of BLAGN2 could reveal rich information about the AGN environment close to the black hole.

In this work, we study the BLAGN in the *XMM-XXL* north survey. We introduce the data in Section 2, investigate the X-ray obscuration of BLAGN in Section 3 and the optical spectral properties of them in Section 4. The results are summarized and discussed in Section 5.

2 THE DATA

2.1 The XXL-BOSS BLAGN sample

The *XMM-XXL* survey provides a large catalogue (8445) of point-like X-ray sources. 3042 of them with *R*-band¹ AB magnitude between 15.0 and 22.5 were followed up by the Baryon Oscillation Spectroscopic Survey (BOSS) spectrograph (Georgakakis & Nandra 2011; Menzel et al. 2016). Based on the widths of the optical emission lines, i.e. $\text{H}\beta$, Mg II , or C IV , Liu et al. (2016) measured the M_{BH} of the BLAGN in the catalogue. For sources with reliable redshift measurement and optical classification, Liu et al. (2016) used a Bayesian method (Buchner et al. 2014) to measure the N_{H} and rest-frame 2–10 keV luminosity L_{X} . To select the X-ray obscured sources, we use the same divide at $N_{\text{H}} = 10^{21.5} \text{ cm}^{-2}$ as used in Merloni et al. (2014), who found that this value provides the most consistent X-ray and optical classifications. Among the XXL BLAGN, >20 per cent of them have $N_{\text{H}} > 10^{21.5} \text{ cm}^{-2}$, and

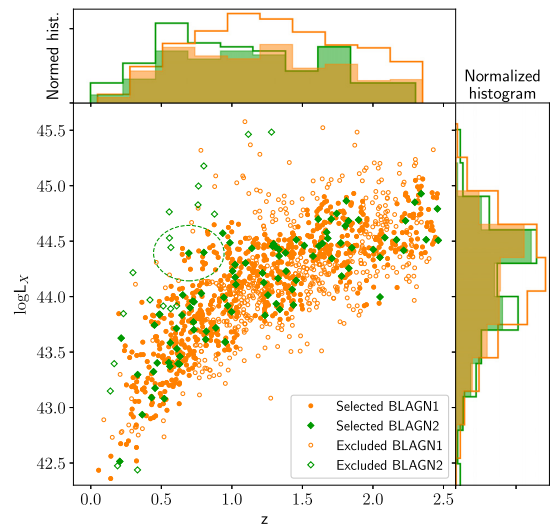


Figure 1. The X-ray luminosity–redshift scatter plot of the BLAGN1 (orange) and BLAGN2 (green) in sample ‘0’. In the ‘same $L_{\text{X}}-z$ ’ reselection (see Section 3.2 for details), the solid points are selected and the empty ones are excluded. The maximum distance of one point is shown with a dashed circle as an example. Normalized histograms of luminosity and redshift are also shown for BLAGN2 (green) and BLAGN1 (orange) samples, in which the empty part corresponds to the excluded sources.

if only sources with >50 net counts are considered, the fraction is ~ 10 per cent.

With respect to the M_{BH} measured on the basis of hydrogen Balmer lines, M_{BH} measured using Mg II is broadly consistent, while C IV -based measurement can be systematically biased (Shen et al. 2008; Shen & Liu 2012; Coatman et al. 2017). Therefore, we select only the sources with M_{BH} measured using $\text{H}\beta$ and Mg II . This is roughly equivalent to excluding the sources at $z \gtrsim 2.5$ whose Mg II line is out of the BOSS wavelength range (3600–10 000 Å). When having both $\text{H}\beta$ and Mg II measurements of M_{BH} , we choose the one with smaller uncertainty.

Based on the X-ray spectral analysis presented in Liu et al. (2016), we classify sources with $\log N_{\text{H}} > 21.5$ and with 1σ lower limit of $\log N_{\text{H}}$ above 21 as BLAGN2, and those with $\log N_{\text{H}} \leq 21.5$ as BLAGN1. We exclude sources with low optical signal-to-noise ratio (S/N ; $\text{SN}_{\text{MEDIAN_ALL}} < 1.6$, see appendix B of Menzel et al. 2016) and a few low X-ray S/N sources whose intrinsic 2–10 keV luminosities are not well constrained (the width of $\log L_{\text{X}}$ 1σ confidence interval > 0.5). We also exclude a few sources whose broad lines are very weak through visual inspection, because they might be actually narrow-line AGN (NLAGN) with false detection of broad lines. By now, our sample comprises 1172 BLAGN1 and 113 BLAGN2, whose luminosity–redshift distributions are shown in the central panel of Fig. 1. A code name ‘0’ is assigned to this sample. However, this is not yet the eventual sample. Thanks to our analysis of the source properties in the following sections, we notice that it is best to further exclude a few sources whose nature is uncertain (highly obscured or having a very low accretion rate). These additional filters give rise to the eventual sample ‘1’, see Section 3.2 for details.

2.2 The BOSS spectra

We show the BOSS spectral shape of our sources in Fig. 2. Although all the sources are defined as BLAGN, whose optical spectral shapes

¹Throughout the paper, *R* band corresponds to the Sloan Digital Sky Survey (SDSS) observed band.

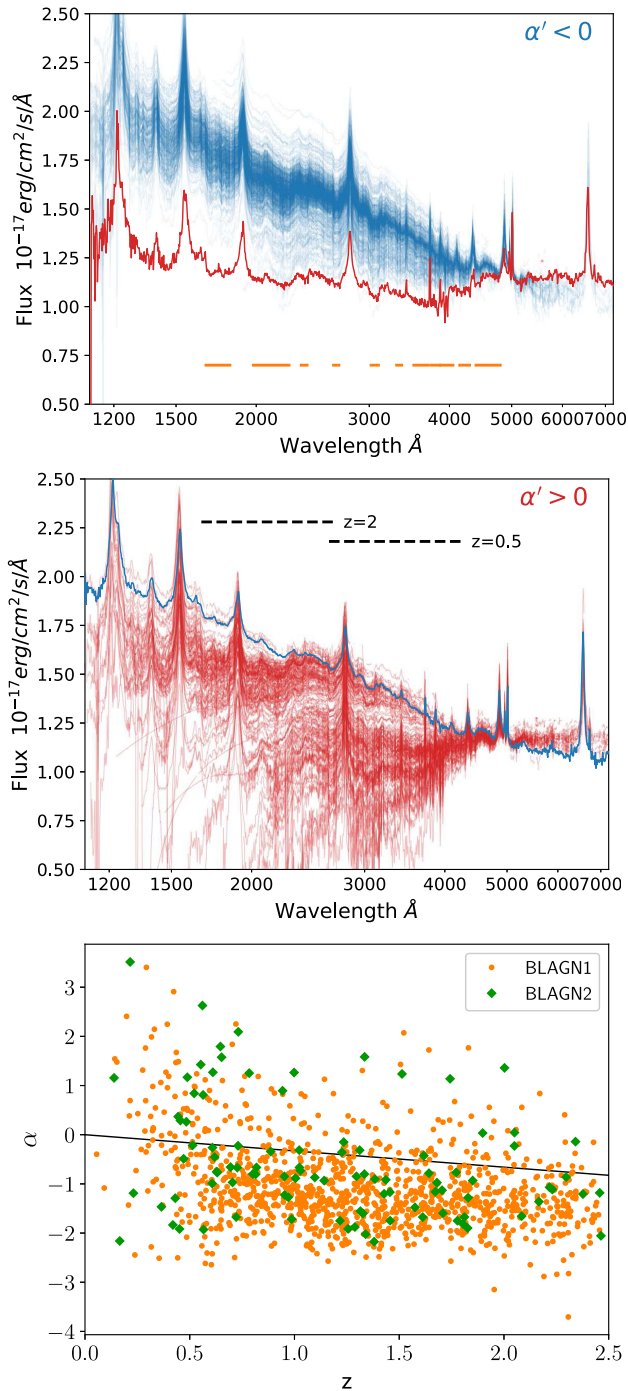


Figure 2. The blue lines in the top panel and the red lines in the middle panel show the Galactic-extinction-corrected, rest-frame spectrum of each source for the blue ($\alpha' < 0$) and red ($\alpha' > 0$) BLAGN, respectively. All the spectra are normalized to the median composite spectrum of the blue BLAGN, which is plotted as the blue line in the middle panel, and the median composite spectrum of the red BLAGN is plotted as the red line in the top panel (see Section 4.2 for details about the composite spectra). We plot the best-fitting models from the SDSS pipeline instead of the real data in order to have a clear look even in the low S/N cases. The ‘line-free’ windows used in the power-law fitting are shown with orange lines in the top panel. Examples of the wavelength spans of the power-law-fitting bands are shown in the middle panel for a redshift of 2 and 0.5. The bottom panel shows the α - z scatter plot of the BLAGN1 (orange) and BLAGN2 (green). The black line ($\alpha = -0.33z$) corresponds to $\alpha' = 0$.

are expected to be a blue power law with a negative slope around -1.5 (Vanden Berk et al. 2001), we find that a fraction of them show continuum reddening. To evaluate the reddening, we define a slope parameter as follows. Since our sources span a wide redshift range, we define the slope on a redshift-dependent rest-frame band. We choose a series of ‘line-free’ sections between rest frame 1670 and 4800 Å,² excluding the line-dominated part but as little as possible, as shown in the top panel of Fig. 2. In these selected sections, for each source, we choose the bluest available part that spans a wavelength width of 0.25 dex (see examples in the middle panel of Fig. 2) to define the slope parameter. After shifting each spectrum to rest frame, we calculate a slope α by a linear fitting in the selected bands.

As shown in the bottom panel of Fig. 2, we find an anticorrelation between α and redshift. This anticorrelation is likely caused by the R -band magnitude selection bias against dust-extinct sources at high redshifts. Because dust extinction affects only the blue band in rest frame, so that the observed R -band flux is less affected for low- z sources than for high- z sources. To skirt around this bias, we define a less-redshift-dependent slope α' by applying a redshift correction $\alpha + 0.33z$ using the slope of the anticorrelation 0.33, as shown in the bottom panel of Fig. 2. Using this α' parameter, we can separate the reddened sources that appear different from the majority of the sample at different redshifts. Hereafter, the sources with $\alpha' > 0$ are called red AGN, and the others are called blue. The top and middle panels of Fig. 2 clearly show the differences between their continuum shapes.

Comparing the median composite spectra (generated using method ‘A’ as described in Section 4.2) between the blue and red BLAGN, we get an $E(B - V) = 0.27$. It corresponds to $A_V = R_V E(B - V) = 1.4$, if we consider the possibility that AGN might have larger dust grain size (e.g. Laor & Draine 1993; Imanishi 2001; Maiolino et al. 2001) and hence adopt $R_V = 5.3$ (Gaskell et al. 2004) as opposed to the Galactic value of 3.1 (but see also Weingartner & Murray 2002; Willott 2005). Note that this value only corresponds to a fraction of low- z sources at $z \lesssim 0.7$. Such low- z sources have significant stellar contamination (see Section 4.4), which could flatten the spectra at $\gtrsim 4000$ Å. Meanwhile, as discussed above, high- z sources show lower extinction because of sample selection effects (see also Willott 2005). Therefore, for the whole red AGN sample, this A_V value is more of a moderate upper limit than a typical value of the optical extinction. Adopting an empirical correlation $N_H = 2.2 \times 10^{21} A_V$ (Güver & Özel 2009), it corresponds to an N_H of $3 \times 10^{21} \text{ cm}^{-2}$ – approximately the lower limit of the X-ray N_H of the BLAGN2. Nevertheless, red AGN only constitute a small fraction of the BLAGN2. In most of the BLAGN2 that are blue (below the black line in the bottom panel of Fig. 2), there is rarely any optical dust extinction. Therefore, we conclude that the dust accountable for the optical extinction is insufficient to explain the X-ray obscuration in the BLAGN2.

3 THE X-RAY OBSCURATION

3.1 The effective Eddington limit

The effective Eddington limit is much lower for dusty gas than for ionized dust-free gas (Laor & Draine 1993; Scoville & Norman 1995; Murray, Quataert & Thompson 2005). Such a limit defines a

²Throughout this paper, the wavelengths correspond to rest frame if not explicitly specified.

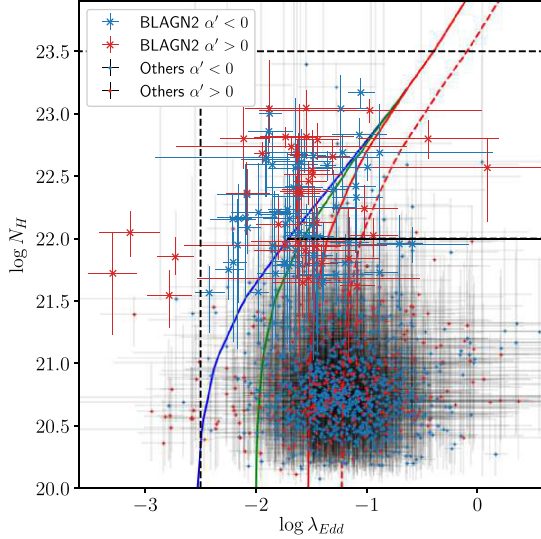


Figure 3. $N_{\text{H}}-\lambda_{\text{Edd}}$ scatter plot of the BLAGN. The crosses (x) indicate BLAGN2, and the plus markers (+) indicate the others, including BLAGN1 and a few low S/N sources. The sources with blue ($\alpha' < 0$) and red ($\alpha' > 0$) optical continua are plotted in blue and red colours, respectively. We show the effective Eddington limit estimated by Fabian et al. (2009) with dust grain abundances of 1 (blue), 0.3 (green), and 0.1 (red). The red dashed line shows a factor of 2 increase in the effective Eddington limit with a grain abundance of 0.1 due to the mass of intervening stars. The black solid line corresponds to the lower boundary of 10^{22} cm^{-2} . The black dashed lines correspond to $N_{\text{H}} = 10^{23.5} \text{ cm}^{-2}$ and $\lambda_{\text{Edd}} = 10^{-2.5}$.

blow-out region (forbidden region) in the $N_{\text{H}}-\lambda_{\text{Edd}}$ plane for AGN, in which the living time of an AGN is expected to be short. In Fig. 3, we plot our sources in the $N_{\text{H}}-\lambda_{\text{Edd}}$ diagram to see if they obey the effective Eddington limit. The λ_{Edd} is calculated from L_{X} assuming a constant bolometric correction factor of 20, as done in Ricci et al. (2017). We show the effective Eddington limits calculated by Fabian et al. (2009) for dusty gas located close to the black hole, where the black hole dominates the mass locally, with dust grain abundances of 1, 0.3, and 0.1. As done in Ricci et al. (2017), we plot a lower boundary of $N_{\text{H}} = 10^{22} \text{ cm}^{-2}$, below which the obscuration might be due to galaxy-scale dust lanes.

Compared with the BLAGN1 at the same λ_{Edd} , a lack of BLAGN2 can be seen in the dust blow-out region at $\lambda_{\text{Edd}} \gtrsim 10^{-1}$, similar as found by Ricci et al. (2017). Using the BAT 70-month AGN catalogue, Ricci et al. (2017) found that among 160 NLAGN with $N_{\text{H}} \geq 10^{22} \text{ cm}^{-2}$, a very small fraction (five sources, 3 per cent) lie in the dust blow-out region (see their fig. 3). Among our BLAGN2 sample, 80 have $N_{\text{H}} > 10^{22} \text{ cm}^{-2}$, and 18 out of the 80 (~ 22 per cent) lie in the dust blow-out region (blue and solid black lines in Fig. 3). The fraction we find is larger. Note that this is not a rigorous comparison, since the BAT survey and XXL survey have very different depths and different selection limits. However, considering that the dust column density revealed by optical extinction is insufficient to account for the X-ray obscuration (Section 2.2), it is likely true that BLAGN2 have a larger probability to occur in the dust blow-out region than NLAGN, in the sense that the X-ray absorbers in BLAGN2 have a lower dust fraction and thus a higher effective Eddington limit. We choose the limit with a grain abundance of 0.1 (red solid line in Fig. 3), and correct it by a factor of 2 to account for the mass of the stars inwards from the obscuring material, as done in Fabian et al. (2009). This factor corresponds to a scale of a few parsec from the nucleus in the case of our Galaxy (Schödel

et al. 2007). Using this corrected limit (red dashed line in Fig. 3), we find that there are only three sources in the blow-out region. Incidentally, using this limit we can efficiently select sources that are likely outflowing (see Appendix B). Therefore, we argue that the X-ray obscuration in BLAGN can be well described by such an absorber with a low dust fraction and located at about a few parsec from the black hole.

3.2 The major difference between BLAGN1 and BLAGN2

In this section, we compare the L_{X} , M_{BH} and λ_{Edd} between the BLAGN1 and BLAGN2, in order to investigate which factor is the main physical driver of the difference between them. First, to reduce sample selection bias and compare one parameter between the BLAGN1 and BLAGN2 with the others under control, we reselect the BLAGN1 and BLAGN2 samples from sample ‘0’.

To compare their M_{BH} at the same L_{X} and z , we select a BLAGN1 sample and a BLAGN2 sample with the same L_{X} and z distribution. We repeatedly select the nearest BLAGN1 to each BLAGN2 in the $\log L_{\text{X}}-z$ space. When a BLAGN2 has no more neighbour found within a maximum distance of 0.25, it is excluded and the procedure is restarted with the reduced BLAGN2 sample. We find that for a subsample of 94 BLAGN2, we could repeat the nearest point selection for eight times; in other words, we could assign eight nearest points from the BLAGN1 to each of these BLAGN2. As shown in Fig. 1, the redshift distributions are significantly different between the BLAGN1 and BLAGN2 in sample ‘0’ (empty histogram), because of the X-ray sample selection bias against BLAGN2 that have relatively lower observed fluxes. In the reselection, the excluded BLAGN2 (empty diamonds) have relatively lower z and higher L_{X} and the excluded BLAGN1 (empty circles) have relatively higher z and lower L_{X} . As a result, the selected BLAGN1 and BLAGN2 have a highly identical $L_{\text{X}}-z$ distribution (filled points and histograms).

To select samples with the same $M_{\text{BH}}-z$ distribution between the BLAGN1 and BLAGN2, we perform a similar sample reselection as above in the $\log M_{\text{BH}}-z$ space, allowing a maximum distance of 0.25. In this selection, the excluded BLAGN2 have relatively lower z and higher M_{BH} and the excluded BLAGN1 have relatively higher z and lower M_{BH} . Similarly, we can also select samples of BLAGN1 and BLAGN2 that have the same $\lambda_{\text{Edd}}-z$ distribution, trimming off a few BLAGN2 with relatively lower z and lower λ_{Edd} and a fraction of BLAGN1 with relatively higher z and higher λ_{Edd} .

There are 11 highly obscured sources with $\log N_{\text{H}} > 23.5$ in the sample ‘0’. We note that all except one of them (with the lowest N_{H}) are excluded in the ‘same $L_{\text{X}}-z$ ’ selection, because they have relatively higher L_{X} than both the BLAGN1 and the other BLAGN2 at the same redshifts. Such high L_{X} of them are possibly overestimations (see Appendix A). To be conservative, we exclude such sources and focus on the others whose N_{H} and L_{X} are better constrained by the *XMM-Newton* spectra. As shown in Fig. 3, most of the sources at λ_{Edd} below $10^{-2.5}$ have a red optical continuum. At such low accretion rates, the sources likely have their optical emission dominated by the host galaxy. We also exclude them from further analysis. Applying the two additional filters ($N_{\text{H}} < 10^{23.5} \text{ cm}^{-2}$ and $\lambda_{\text{Edd}} > 10^{-2.5}$), as shown by the black dashed line in Fig. 3, we select a sample ‘1’ from sample ‘0’. All the analyses below are based on sample ‘1’.

Performing the sample reselections described above on sample ‘1’, we select three pairs of samples:

‘= L_{X} ’: 92 BLAGN2 and 7×92 BLAGN1 with the same $L_{\text{X}}-z$ distribution;

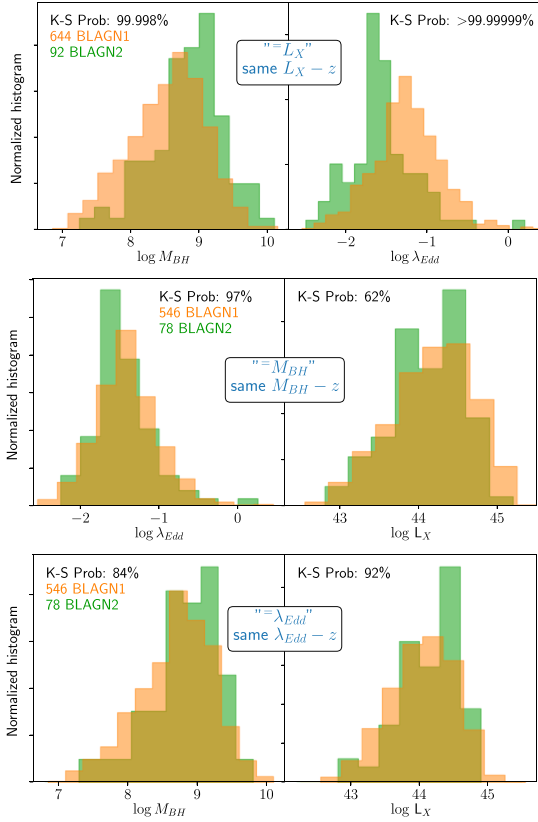


Figure 4. Comparison of distributions of M_{BH} , L_X , or λ_{Edd} between the BLAGN1 and BLAGN2 using a series of samples – ‘ L_X ’, ‘ M_{BH} ’, and ‘ λ_{Edd} ’, from top to bottom. Number of sources and K–S test probability are shown in each panel.

‘ M_{BH} ’: 78 BLAGN2 and 7×78 BLAGN1 with the same $M_{\text{BH}}-z$ distribution;

‘ λ_{Edd} ’: 78 BLAGN2 and 7×78 BLAGN1 with the same $\lambda_{\text{Edd}}-z$ distribution.

We compare the M_{BH} and λ_{Edd} between the ‘ L_X ’ BLAGN1 and BLAGN2 samples, compare the λ_{Edd} and L_X between the ‘ M_{BH} ’ samples, and compare the M_{BH} and L_X between the ‘ λ_{Edd} ’ samples. For each comparison we perform a Kolmogorov–Smirnov (K–S) test.³

As shown in Fig. 4, at the same L_X (top panel, ‘ L_X ’), the BLAGN2 have significantly higher M_{BH} and lower λ_{Edd} . The median M_{BH} and λ_{Edd} of the ‘ L_X ’ samples are 8.60, -1.26 for the BLAGN1 and 8.93, -1.60 for the BLAGN2. At the same M_{BH} (middle panel, ‘ M_{BH} ’), the L_X is similar between the BLAGN1 and BLAGN2; the λ_{Edd} is lower in the BLAGN2 but only slightly (about 2σ). At the same λ_{Edd} (bottom panel, ‘ λ_{Edd} ’), we find no significant differences either on the other two parameters between the BLAGN1 and BLAGN2. Therefore, among L_X , M_{BH} , and λ_{Edd} , the main physical difference between the BLAGN1 and BLAGN2 consists in M_{BH} or λ_{Edd} . However, as will be discussed in Section 5.4, the M_{BH} difference can be alternatively attributed to a bias caused by an inclination effect.

A constant bolometric correction factor is used to calculate λ_{Edd} from L_X (Section 3.1). However, if we consider the correlation

³Throughout this paper, the K–S test probability denotes 1 minus the probability that two samples are drawn from the same population.

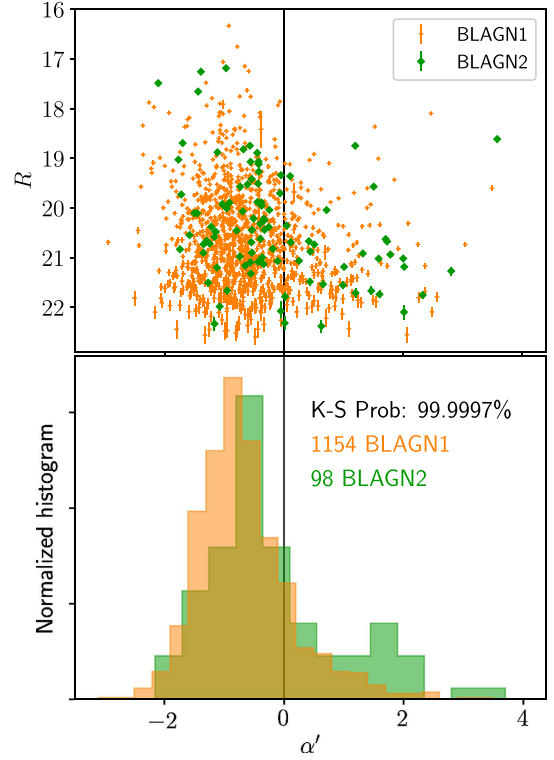


Figure 5. The scatter plot of R -band magnitude R and α' (upper panel) and the normalized α' distributions (lower panel) of the BLAGN1 (orange) and BLAGN2 (green) of sample ‘1’. The vertical black line corresponds to $\alpha' = 0$. The error bars of R correspond to 1σ .

between the X-ray bolometric correction factor and λ_{Edd} (Vasudevan & Fabian 2007, 2009; Lusso et al. 2012; Liu et al. 2016), the λ_{Edd} of BLAGN2 can be even lower and thus the difference between the λ_{Edd} of the BLAGN1 and BLAGN2 can be even stronger.

4 THE OPTICAL SPECTRA

4.1 The optical extinction

In Section 2.2, we have defined a slope parameter α' to evaluate the optical continuum reddening. In the upper panel of Fig. 5, we plot α' against the observed R -band magnitude for sample ‘1’. Clearly, the red ($\alpha' > 0$) sources have relatively lower optical fluxes, indicating that the reddening should be mainly caused by dust extinction, which reduces the optical flux.

As a first, rough comparison between the optical continuum of the BLAGN1 and BLAGN2, we compare their slopes α' using the sample ‘1’ in the lower panel of Fig. 5. We find that the slopes are significantly redder in the BLAGN2 sample at a K–S test confidence probability of 99.9997 per cent. Using the ‘ L_X ’, ‘ M_{BH} ’, or ‘ λ_{Edd} ’ samples, we also find such differences at high K–S test probabilities of 99.9992, 99.8, and 99.94 per cent, respectively. Therefore, the optical dust extinction is correlated with the X-ray obscuration, and this correlation is not driven by the λ_{Edd} or M_{BH} difference between the BLAGN1 and BLAGN2 as we find in Section 3.2.

With stronger dust extinction, the BLAGN2 can be more easily missed by the R -band magnitude-limited sample selection threshold. Taking account of this bias will boost the difference of dust extinction levels between the BLAGN1 and BLAGN2.

4.2 Spectral stacking methods

In order to compare the optical spectra of the BLAGN1 and BLAGN2 in more details, we stack their SDSS spectra. Three different methods of stacking, A, B, and C, will be used below for different purposes. In method ‘A’, we calculate the geometric mean or median of the normalized spectrum of each source in order to study the continuum shape. First, for each spectrum, we apply Galactic extinction correction using the extinction function of Cardelli, Clayton & Mathis (1989) and exclude the bins with high sky background or with observed wavelength below 3700 or above 9500 Å. Then, we shift the spectrum to rest frame using a bin size of $1e - 4$ dex, which is the same as the bin size of the original spectrum, and select the 100 available bins with the longest wavelength below rest frame 5100 Å as the normalizing window for each spectrum. All the spectra are then ordered by redshift, and, starting from the second one, each spectrum is normalized in the selected window to the composite spectrum of the sources with lower redshifts. To calculate the normalization factor, we use the best-fitting models from the SDSS pipeline instead of the real spectra to avoid the high variance of the spectra in some low S/N cases (see Fig. 2 for an illustration of the normalization). Using the extinction-corrected, shifted, and normalized spectra, a composite spectrum (geometric mean or median) is calculated. Then we repeat the normalizing procedure but normalizing each spectrum to the generated composite spectrum instead of normalizing each spectrum to the ones with lower redshifts. The composite spectrum converges after a few iterations, then the 68 per cent confidence intervals are measured using the bootstrap percentile method. Using this method, we present the composite spectrum of our BLAGN in Appendix C.

Stacking method ‘B’ is used to compare the optical fluxes between samples. In this method, we calculate the median of all the sources in each wavelength bin without normalizing the spectra to each other. Instead, each extinction-corrected and shifted spectrum is multiplied by $D_L^2(1+z)$, where D_L is the luminosity distance, in order to preserve the luminosity. The 16 and 84 per cent percentile spectra are used to estimate the flux scatter. Note that such a composite spectrum does not represent the spectral shapes of the sources, it shows the fluxes of the sources at specific redshifts – low- z sources dominate the red part and high- z sources dominate the blue part.

To measure the line features in the composite spectra, we use a stacking method ‘C’, which calculates the median spectra of the ratios of each spectrum to its local best-fitting continuum. For two sets of lines – [Ne III]–Ca II and H β –[O III] – we fit the local continuum of each source in two different sets of bands – 3750–3800, 3885–3910, 3947–3960, and 3990–4050 for the former, 4600–4750 and 5050–5200 for the latter. Rather than measuring line equivalent width (EW) accurately, we aim at making a comparison of the line EW between different types of sources. Therefore, we just fit the spectra with a simple power law in the selected windows and calculate the data to model ratio. The median of the ratios are calculated as the composite line spectrum.

When studying the continuum with methods ‘A’ and ‘B’ as above, we bin the spectra by a factor of 8 to reduce fluctuation. To study the line features, we use the unbinned spectra with the original bin size of $1e - 4$ dex. Sources at different redshifts contribute to the composite line spectra for different line sets. If a source has less than 10 available bins on the left- or right-hand side of the Ca II or

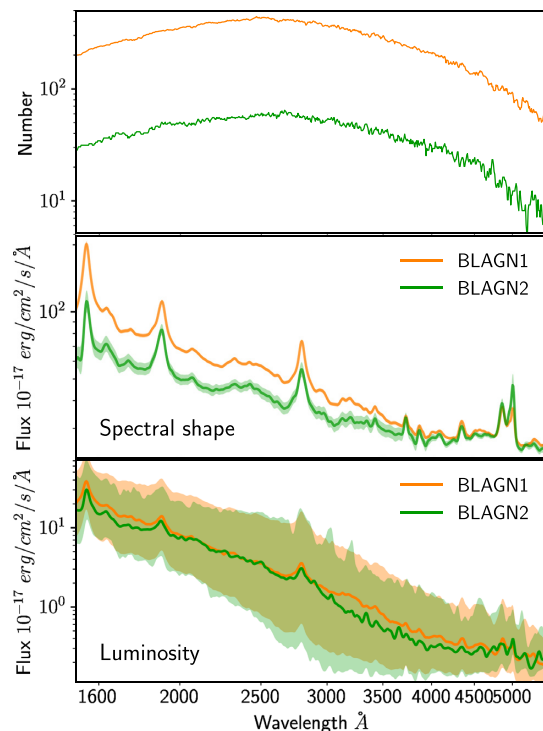


Figure 6. The top panel shows the number of contributing sources per bin. The middle panel shows the geometric mean spectra and 1σ error of the BLAGN1 (orange) and BLAGN2 (green) generated using method ‘A’, which represent the mean spectral shapes. The bottom panel shows the composite spectra generated using method ‘B’ and the 68 per cent scatter of the samples, which represent the mean luminosities. The ‘ $=M_{\text{BH}}$ ’ samples are used.

[O III] line, it is excluded from the stacking of the corresponding line set.

4.3 The optical continuum

To compare the optical continuum shape between the BLAGN1 and BLAGN2, we stack the ‘ $=M_{\text{BH}}$ ’ BLAGN1 and BLAGN2 spectra, respectively, using method ‘A’. As shown in Fig. 6, having the same redshift distribution, the curves of source number per bin have the same shape for the BLAGN1 and BLAGN2 samples (the top panel). The composite spectrum of BLAGN2 is much flatter than that of BLAGN1 (the middle panel), showing the higher probability of continuum reddening in BLAGN2.

In order to check whether the optical luminosities of BLAGN2 are reduced by dust extinction, we compare the composite spectra generated using method ‘B’ between the ‘ $=M_{\text{BH}}$ ’ BLAGN1 and BLAGN2 samples. As shown in the bottom panel of Fig. 6, we find no significant difference. It is because the dust extinction occurs in both BLAGN1 and BLAGN2 and in only a small fraction of them. Taking also the large luminosity scatter into account, the dust extinction could not reduce the mean luminosity of the whole BLAGN1 or BLAGN2 sample significantly.

Comparing the composite spectra between the BLAGN1 and BLAGN2 using the ‘ $=L_X$ ’ or ‘ $=\lambda_{\text{Edd}}$ ’ samples, we find similar results about both the spectral shape and the optical luminosity. As discussed in Section 4, the more severe continuum reddening in BLAGN2 than in BLAGN1 is not driven by differences in λ_{Edd} or M_{BH} ; it is just associated with the X-ray obscuration.

4.4 The line features

We have seen in the previous sections that the X-ray obscuration in BLAGN2 is statistically associated with a higher level of optical spectral reddening, which is likely caused by dust extinction. In order to further understand the reason of the reddening, we compare the line features, not only between the BLAGN1 and BLAGN2, but also between the red ($\alpha' > 0$) and blue ($\alpha' < 0$) BLAGN. This is because the BLAGN1 and BLAGN2 samples have highly overlapped dust reddening distributions (Fig. 5), but the red and blue BLAGN samples separate the sources with relatively low and high reddening levels distinctly. Note that, at low redshifts, the red BLAGN sample also tends to select sources with strong stellar contaminations in the optical spectra. In this section, we divide the sample ‘1’ into red and blue subsamples and into BLAGN1 and BLAGN2 samples, and then make the composite line spectra for two sets of lines – [Ne III]–Ca II and H β –[O III] – using method ‘C’ (see Section 4.2) for each of the four samples. Using the ‘= L_X ’, ‘= M_{BH} ’, or ‘= λ_{Edd} ’ samples instead does not change the results obtained in this section, because the optical spectral difference between the BLAGN1 and BLAGN2 is not driven by any physical parameters (L_X , M_{BH} , or λ_{Edd}), as noticed previously in Sections 4.1 and 4.3. The composite line spectra are shown in Fig. 7.

To estimate the line EW, we fit the narrow [Ne III] and Ca II lines with single-Gaussian profiles and fit the H β and [O III] 5007 lines with double-Gaussian profiles (see Fig. 7). The line EW can be affected by two factors in opposite ways: dust extinction enhances line EW by reducing the underlying continuum and stellar contamination reduces line EW by enhancing the underlying continuum. For each pair of lines, the two lines have similar wavelengths, so that the impact of dust extinction should be similar. Suppose the local continuum of the blue BLAGN (BLAGN1) is composed of a power-law component p and a galaxy component g , and in the red BLAGN (BLAGN2), the power-law emission is reduced by a factor of $1 - a$ ($a < 1$), the galaxy emission is increased by a factor of $1 + b$, and the narrow-line flux remains the same. With respect to the blue BLAGN (BLAGN1), the EW of AGN emission line in the red BLAGN (BLAGN2) is enhanced by a factor of $\frac{p+g}{p-ap+g+bg}$, and the EW of stellar absorption line is enhanced by a factor of $\frac{(p+g)(1+b)}{p-ap+g+bg}$.

The first pair of lines to consider are the AGN emission line [Ne III] 3869 and the galaxy absorption line Ca II K 3934 (the upper two panels of Fig. 7). We find that, compared with the blue BLAGN (BLAGN1), both lines are enhanced in the red BLAGN (BLAGN2). It indicates that the impact of dust extinction is stronger than that of stellar contamination ($ap > bg$). In the comparison between the blue and red BLAGN, the enhancement amplitude of the stellar absorption line (Ca II) is much larger than that of the AGN emission line ([Ne III]), indicating that the stellar contamination is strong ($1 + b$ is significantly > 1). It is not the case in the comparison between the BLAGN1 and BLAGN2, indicating that the stellar components are similar between them ($b \ll 1$).

The second pair of lines to consider are the broad H β line and the narrow [O III] 5007 line. We find that they are enhanced in the BLAGN2 with respect to the BLAGN1, but not in the red BLAGN with respect to the blue BLAGN. It is because at such long wavelengths (~ 5000 Å), the relative strength of the stellar component (g/p) is much larger than at ~ 4000 Å, so that in the latter case (blue versus red BLAGN) stellar contamination effect becomes strong enough to counteract the dust extinction effect ($bg \gtrsim ap$). In the former case, the stellar contamination does not make a significant difference between the BLAGN1 and BLAGN2 ($b \ll 1$) and the

major difference consists in the dust extinction. The impact of extinction at the [O III] wavelength can be stronger than at ~ 4000 Å, because the shorter wavelength section corresponds to sources at higher redshifts, where the sample selection biases against sources with high extinction levels (see the bottom panel of Fig. 2).

In the meanwhile, we find that the relative strength (EW ratio) of the H β broad line with respect to the [O III] narrow lines is weaker in the red BLAGN than in the blue ones and also weaker in the BLAGN2 than in the BLAGN1. In other words, at the same [O III] luminosity, the broad H β line luminosity is lower when dust extinction occurs (see Appendix B for examples). It indicates that the optical absorber of the accretion disc could partially block the BLR.

To test the possibility of partially blocked BLR, we calculate the relative strength of the [O III], H β , and Mg II lines with respect to X-ray as the deviation of the $\log L_{\text{line}}$ from the best-fitting line of $\log L_{\text{line}} - \log L_X$, using the L_{line} from the DR9 quasar catalogue built by Shen et al. (2011) as an extension of the DR7 catalogue. The relative line strength is compared between the blue and red BLAGN in Fig. 8. We find that, for the narrow [O III] line, the relative strength is similar between the blue and red AGN. For the broad H β and Mg II lines, the relative strength is significantly weaker in the red AGN than in the blue ones. A similar comparison between the BLAGN1 and BLAGN2 do not show any significant difference because of a few reasons – the highly overlapped extinction levels of the BLAGN1 and BLAGN2 (Fig. 5), the small sample size of the BLAGN2, and the extra scatter introduced by the L_X .

We conclude that, in some BLAGN, the optical absorber could partially block the BLR. In this sense, the dust extinction in BLAGN is similar to the case of NLAGN, where an absorber at a scale between the BLR and the NLR blocks the former and not the latter. As shown in Fig. 7, the H β EW is larger in BLAGN2 than in BLAGN1, because the higher dust extinction in BLAGN2 reduces the continuum more significantly than the broad line.

The M_{BH} of the BLAGN is calculated on the basis of the broad-line full width at half-maximum (FWHM) and optical continuum luminosity. Practically, the continuum luminosity is substituted with broad-line luminosity (Shen et al. 2011). Therefore, the partial covering of BLR in the BLAGN2 could cause an underestimation of their M_{BH} . This bias cannot be strong, considering that the difference of the relative broad line strength between the BLAGN1 and BLAGN2 can only be revealed in terms of the ratio of the median H β EW to the median [O III] EW (Fig. 7) but not in terms of the relative broad-line luminosity to X-ray luminosity. However, taking this into account will boost the difference we find between the M_{BH} or λ_{Edd} of the BLAGN1 and BLAGN2.

5 CONCLUSION AND DISCUSSION

On the basis of the *XMM-Newton* X-ray spectra analysis in the *XMM-XXL* survey and the optical spectroscopic follow-up of the XXL sources in the SDSS-BOSS survey, we compare the BLAGN1 and BLAGN2 to study their X-ray obscuration and related properties. The results are summarized and explained as follows.

5.1 The X-ray absorber

We find that, at the same L_X , BLAGN2 have significantly higher M_{BH} and lower λ_{Edd} than BLAGN1; while at the same M_{BH} or λ_{Edd} , no significant difference about L_X is found between BLAGN1 and BLAGN2 (Fig. 4). In other words, the major difference between BLAGN1 and BLAGN2 consists in M_{BH} or λ_{Edd} and not in L_X .

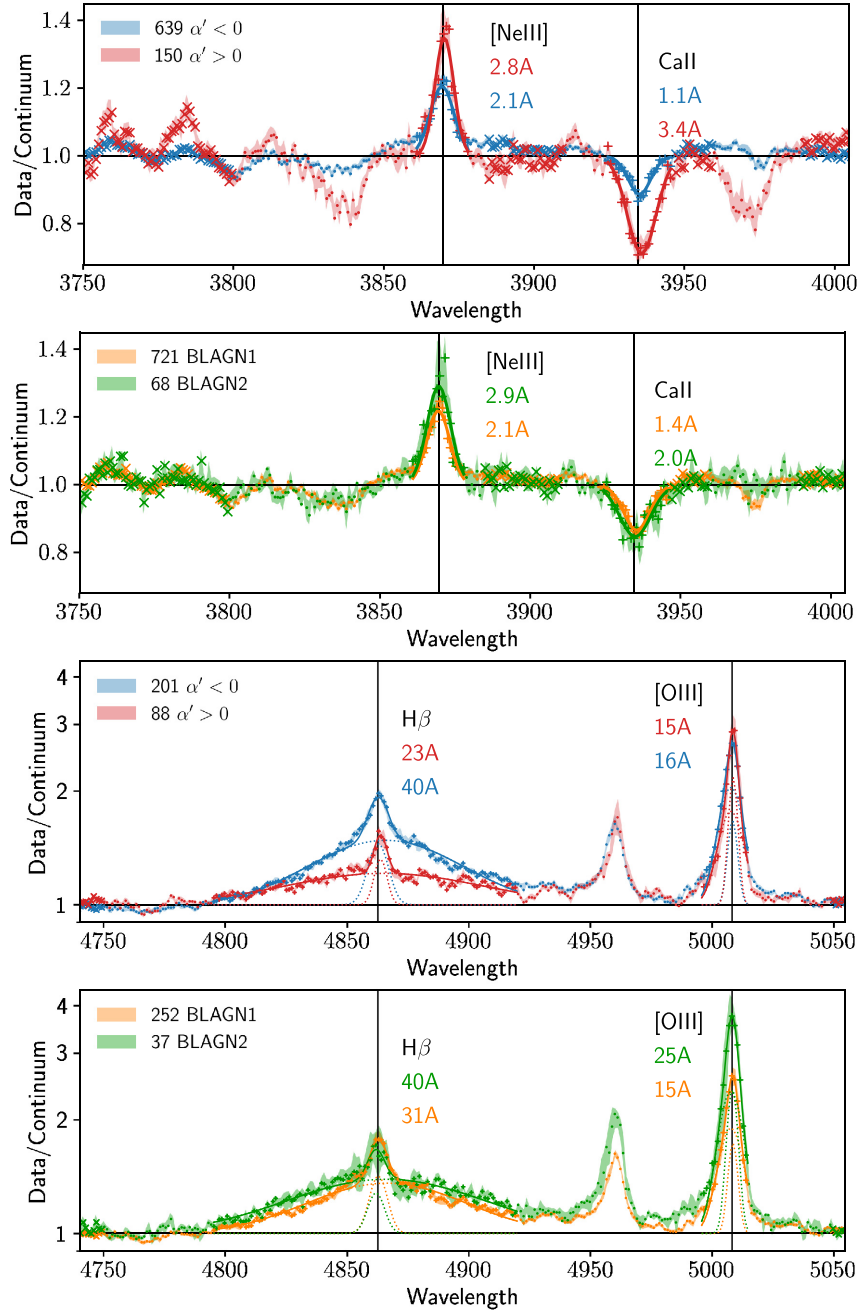


Figure 7. Comparison of the composite line spectra between the blue ($\alpha' < 0$, blue colour) and red ($\alpha' > 0$, red colour) AGN and between the BLAGN1 (orange) and BLAGN2 (green), using sample ‘1’. The upper two panels show the Ne III 3869 emission line and the Ca II K 3934 absorption line. The lower two panels show the broad H β and narrow O III lines. The EW of the lines is marked in the figure using corresponding colours. The best-fitting profiles (single or double Gaussian) are plotted with solid lines. In the cases of double-Gaussian models (for H β and O III), each individual Gaussian profile is plotted with dotted lines. The data points involved in the line fitting are marked with plus crosses. The points involved in the local continuum fitting are marked with crosses (x). The vertical lines correspond to the rest-frame wavelengths of the lines.

In the space of $N_{\text{H}}-\lambda_{\text{Edd}}$, we find a significant lack of BLAGN2 above the effective Eddington limit of a low dust fraction, where the absorber can be swept out by radiation pressure. These properties of the X-ray absorbers in BLAGN are similar as those in NLAGN (Ricci et al. 2017).

One possibility to explain the X-ray obscuration in BLAGN is to cut off the relation between the non-simultaneous X-ray and optical observations by invoking a small X-ray obscuring cloud, which has moved away during the optical follow-up or being too small

to block the extended optical emitting region (disc and BLR) ever. However, the significant difference of M_{BH} and λ_{Edd} between the BLAGN1 and BLAGN2 indicates that such a possibility could only be a minor factor and there should be an intrinsic difference between them.

Unlike the optically thick dust component in NLAGN, whose column density is too high to be measured by means of transmitted optical emission, the optical dust extinction in BLAGN is thin and occasional. Such a thin dusty absorber is far from enough to

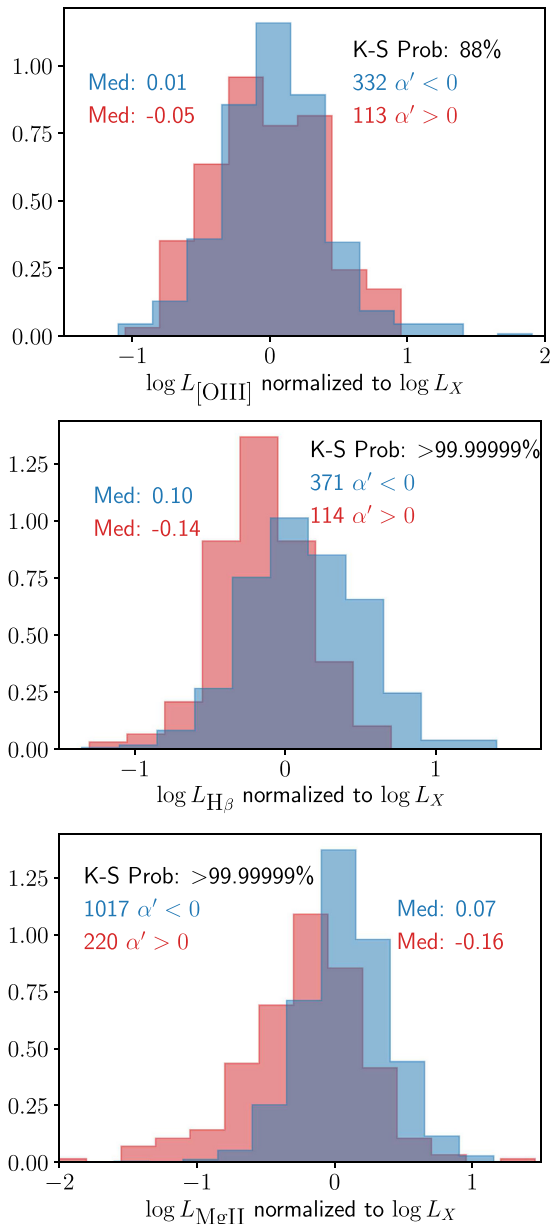


Figure 8. Comparisons of the line luminosities normalized to L_X for [O III], $H\beta$, and Mg II between the blue ($\alpha' < 0$) and red ($\alpha' > 0$) BLAGN using sample ‘1’. We mark the sample sizes, median values, and K–S test probabilities in the figure.

explain the X-ray absorption in the BLAGN2 (Section 2.2). The line-of-sight absorbers in the BLAGN2 must have a low overall dust fraction (Section 3.1), either in terms of a low dust-to-gas ratio, or in terms of a multilayer absorber composed of an inner gas component and an outer dust component. Meanwhile, as revealed by infrared (IR) emission, the dust column densities in NLAGN also appear lower than the X-ray obscuring column density (e.g. Alonso-Herrero, Ward & Kotilainen 1997; Granato, Danese & Franceschini 1997; Fadda et al. 1998; Georgantopoulos et al. 2011; Burtscher et al. 2016). Therefore, in both NLAGN and BLAGN2, the X-ray absorption is at least partially due to a dust-free component.

5.2 The optical absorber

A small fraction of the BLAGN show optical continuum reddening caused by dust extinction. The reddening occurs in both BLAGN1 and BLAGN2, however, BLAGN2 have a higher probability to be reddened than BLAGN1 (Fig. 5), giving rise to a flatter composite spectrum of BLAGN2 than that of BLAGN1 (Fig. 6).

The median EW of a few optical line features, as measured through composite spectra, is compared between the optical red ($\alpha' > 0$) and blue ($\alpha' < 0$) BLAGN and between the BLAGN1 and BLAGN2 (Fig. 7). We find that, in the case of red versus blue BLAGN, both dust extinction and stellar contamination affect the line EW. In the case of BLAGN1 versus BLAGN2, stellar contamination does not make a significant difference between them and the major difference consists in the higher dust extinction level in the BLAGN2, which enhances the line EW in BLAGN2 with respect to BLAGN1.

Using the median EW of the broad $H\beta$ line and the narrow [O III] line, we find that the relative strength of $H\beta$ with respect to [O III] is weaker in the red AGN (BLAGN2) than in the blue AGN (BLAGN1) (Fig. 7). We also find that the relative strength of the broad $H\beta$ and Mg II line luminosities with respect to the X-ray luminosities is weaker in the red AGN than in the blue AGN (Fig. 8). They indicate a partial-covering obscuration on the BLR.

To summarize, the X-ray obscuration in BLAGN tends to coincide with optical dust extinction, which is optically thinner than that in the NLAGN and can be partial covering to the BLR.

5.3 A geometrical torus model

We summarize the properties of the obscuring material in BLAGN as follows.

- (i) The X-ray absorber in BLAGN2 is similar as in NLAGN but has an optically thinner dust component.
- (ii) The accretion disc in BLAGN2 suffers more dust extinction than in BLAGN1 but, of course, not as thick as in NLAGN.
- (iii) In dust-extinct BLAGN, the BLR could also be dust extinct, similarly as in NLAGN, but by a partial covering and optically thinner absorber.

Clearly, from both the X-ray and optical point of view, BLAGN2 take an intermediate place between BLAGN1 and NLAGN. As described below, such an intermediate type can be naturally explained by a multicomponent, clumpy torus model.

It was known from the very beginning that torus is most likely clumpy (Krolik & Begelman 1988). Observationally, clumpy torus models are supported by the fast X-ray absorption variability (e.g. Kaastra et al. 2014; Markowitz, Krump & Nikutta 2014; Marinucci et al. 2016) and the isotropy level of IR emission (e.g. Levenson et al. 2009; Ramos Almeida et al. 2011). They have also been successful in explaining the spectral energy distribution (SED) and IR spectroscopy of AGN (Mor et al. 2009; Ramos Almeida et al. 2009; Hönig et al. 2010; Alonso-Herrero et al. 2011; Lira et al. 2013). We illustrate a clumpy torus in Fig. 9. The dashed-line circle around the central engine (disc+corona) indicates the sublimation radius. The classical dusty torus (blue filled hexagons) is located outside this radius. However, if defined as the X-ray absorber, the torus should extend into this radius and have a gaseous part (green empty circles). This part might contribute in or be identical to the BLR (Goad, Korista & Ruff 2012; Davies et al. 2015). In some local AGN, IR interferometric observations find an additional dust component in the polar region (blue empty hexagons) besides the

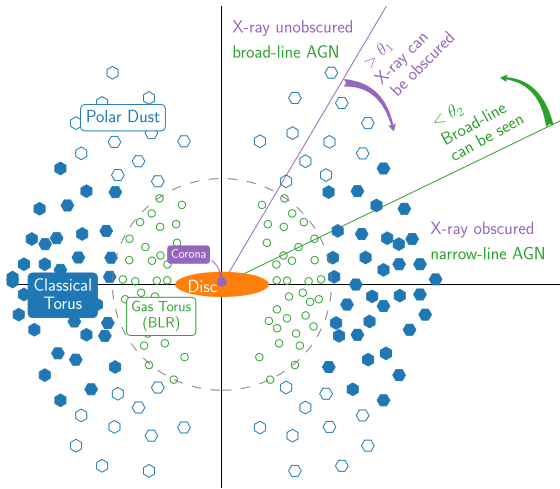


Figure 9. Cartoon of the circumnuclear environment of AGN (not to scale). The dashed-line circle indicates the sublimation radius. Circular clouds inside the sublimation radius are dust free, and hexagon clouds outside the radius are dusty. The empty clouds (BLR and polar dust) are optically thin to the optical emission, and the filled ones are optically thick. Above an inclination angle θ_1 , the overall column density of the gas and dust becomes sufficient to obscure the X-ray emission from the corona (purple). Below an inclination angle θ_2 , the BLR could leak through the optically thick dusty torus.

classical torus. This component is optically thin but emits efficiently in mid-IR, possibly due to an outflowing dusty wind or due to dust in the near-IR (Hönig et al. 2012, 2013; López-Gonzaga et al. 2014, 2016; Tristram et al. 2014; Asmus, Hönig & Gandhi 2016). All the three components – the classical dusty torus, the gaseous inner torus (or BLR), and the polar dust – contribute in the X-ray obscuration. However, the last one can be negligible in terms of column density compared with the other two.

Under the clumpy torus model, the incidence of obscuration along the line of sight is probabilistic in nature. However, considering the geometric structure of three obscuring components as shown in Fig. 9, the obscuring possibility clearly increases with the inclination angle. We can imagine an inclination angle θ_1 , above which the corona becomes X-ray obscured, and an inclination angle θ_2 , below which the BLR can be seen. The typical BLAGN1 and NLAGN are seen at low inclination angles $<\theta_1$ and at high inclination angles $>\theta_2$, respectively. Among the three components, only the equatorial dusty clouds (blue filled hexagons) could efficiently block the BLR; the optically thin polar dust (blue empty hexagons) might only reduce the broad line flux moderately. Also considering that the BLR is an extended region with a much larger scale than that of the corona, it is natural that $\theta_2 > \theta_1$. The intermediate-inclination region between θ_1 and θ_2 is where the BLAGN2 reside. As illustrated in Fig. 9, the optical extinction of the BLAGN2 can be attributed to either the optically thin polar dust or the dusty clumps of the classical torus. In the latter case, the dust extinction can be optically thin in terms of a small line-of-sight number of clumps. Future multiband spectroscopic surveys might allow us to constrain the model in quantity by means of the fraction of BLAGN2 among the entire AGN population.

5.4 Physical driver of the obscuration incidence

We have shown in Section 3.2 that BLAGN2 have higher single-epoch M_{BH} and thus lower λ_{Edd} than BLAGN1. It is possible that the

main physical driver of whether the X-ray emission of a BLAGN is obscured is the λ_{Edd} , which regulates the covering factor of the X-ray absorber by means of radiation pressure, as pointed out by Ricci et al. (2017) for the X-ray obscuration in NLAGN. However, in the framework of the torus model described above, we can alternatively attribute all the differences between BLAGN1 and BLAGN2 to an inclination effect without invoking the λ_{Edd} -driven effect.

We notice that the higher M_{BH} of our BLAGN2 is entirely caused by their larger FWHM of broad lines. It has been shown by plenty of works that broad-line FWHM increases with increasing inclination angle, likely because of a disc-like structure of BLR, and the virial f factor should decrease with increasing inclination angle (Wills & Browne 1986; Risaliti et al. 2011; Pancoast et al. 2014; Shen & Ho 2014; Bisogni et al. 2017b; Mejía-Restrepo et al. 2018). As discussed in Section 5.3, BLAGN2 can be explained as BLAGN with high inclination angles. As a consequence, the larger M_{BH} of BLAGN2 could just result from the failure to consider the inclination effect in the M_{BH} calculation.

Obviously, there is a degeneracy between the λ_{Edd} -driven effect and inclination effect in explaining the incidence of obscuration. These two explanations are not mutually exclusive. However, we remark that, in the framework of our multicomponent, clumpy torus model, the inclination effect simultaneously explains all the findings of this work, including the existence of BLAGN2, the correlation between M_{BH} and X-ray obscuration, the correlation between X-ray obscuration and optical extinction, and the correlation between relative broad line strength and optical extinction. Meanwhile, attributing the larger M_{BH} of BLAGN2 to larger inclination angles, our model naturally explains why we find a correlation between the M_{BH} and X-ray obscuration but not between the M_{BH} and optical extinction. As illustrated in Fig. 9, the X-ray absorber, composed of the BLR (inner gaseous torus) and the classical dusty torus, has a toroidal or disc-like shape. It is strongly anisotropic even within the inclination range of BLAGN ($<\theta_2$), presenting a steep gradient of the average X-ray obscuring column density as a function of inclination angle. However, the optical absorber of BLAGN, composed of the polar dust and the low-inclination, low-density part of the classical dusty torus, is more evenly distributed within the BLAGN inclination range. The anisotropy of the optical absorber could become prominent only when it comes into the regime of NLAGN.

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REFERENCES

- Alonso-Herrero A., Ward M. J., Kotilainen J. K., 1997, *MNRAS*, 288, 977
 Alonso-Herrero A. et al., 2011, *ApJ*, 736, 82
 Antonucci R., 1993, *ARA&A*, 31, 473
 Asmus D., Hönic S. F., Gandhi P., 2016, *ApJ*, 822, 109
 Bisogni S., di Serego Alighieri S., Goldoni P., Ho L. C., Marconi A., Ponti G., Risaliti G., 2017a, *A&A*, 603, A1
 Bisogni S., Marconi A., Risaliti G., 2017b, *MNRAS*, 464, 385
 Boroson T. A., Green R. F., 1992, *ApJS*, 80, 109
 Brightman M., Nandra K., 2011, *MNRAS*, 413, 1206
 Brightman M., Nandra K., Salvato M., Hsu L.-T., Aird J., Rangel C., 2014, *MNRAS*, 443, 1999
 Brusa M. et al., 2003, *A&A*, 409, 65
 Brusa M. et al., 2015, *MNRAS*, 446, 2394
 Buchner J. et al., 2014, *A&A*, 564, A125
 Burlon D., Ajello M., Greiner J., Comastri A., Merloni A., Gehrels N., 2011, *ApJ*, 728, 58
 Burtscher L. et al., 2016, *A&A*, 586, A28
 Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
 Coatman L., Hewett P. C., Banerji M., Richards G. T., Hennawi J. F., Prochaska J. X., 2017, *MNRAS*, 465, 2120
 Davies R. I. et al., 2015, *ApJ*, 806, 127
 Fabian A. C., Vasudevan R. V., Mushotzky R. F., Winter L. M., Reynolds C. S., 2009, *MNRAS*, 394, L89
 Fadda D., Giuricin G., Granato G. L., Vecchies D., 1998, *ApJ*, 496, 117
 García-Burillo S. et al., 2016, *ApJ*, 823, L12
 Gaskell C. M., Goosmann R. W., Antonucci R. R. J., Whysong D. H., 2004, *ApJ*, 616, 147
 Georgakakis A., Nandra K., 2011, *MNRAS*, 414, 992
 Georgantopoulos I. et al., 2011, *A&A*, 531, A116
 Goad M. R., Korista K. T., Ruff A. J., 2012, *MNRAS*, 426, 3086
 Granato G. L., Danese L., Franceschini A., 1997, *ApJ*, 486, 147
 Güver T., Özel F., 2009, *MNRAS*, 400, 2050
 Harrison C. M., Alexander D. M., Mullaney J. R., Swinbank A. M., 2014, *MNRAS*, 441, 3306
 Hasinger G., 2008, *A&A*, 490, 905
 Ho L. C., Goldoni P., Dong X.-B., Greene J. E., Ponti G., 2012, *ApJ*, 754, 11
 Hönic S. F., Kishimoto M., Gandhi P., Smette A., Asmus D., Duschl W., Polletta M., Weigelt G., 2010, *A&A*, 515, A23
 Hönic S. F., Kishimoto M., Antonucci R., Marconi A., Prieto M. A., Tristram K., Weigelt G., 2012, *ApJ*, 755, 149
 Hönic S. F. et al., 2013, *ApJ*, 771, 87
 Imanishi M., 2001, *AJ*, 121, 1927
 Imanishi M., Nakanishi K., Izumi T., Wada K., 2018, *ApJ*, 853, L25
 Kaastra J. S. et al., 2014, *Science*, 345, 64
 Kakkad D. et al., 2016, *A&A*, 592, A148
 Koss M. et al., 2017, *ApJ*, 850, 74
 Krolik J. H., Begelman M. C., 1988, *ApJ*, 329, 702
 Laor A., Draine B. T., 1993, *ApJ*, 402, 441
 Lawrence A., Elvis M., 1982, *ApJ*, 256, 410
 Levenson N. A., Radomski J. T., Packham C., Mason R. E., Schaefer J. J., Telesco C. M., 2009, *ApJ*, 703, 390
 Lira P., Videla L., Wu Y., Alonso-Herrero A., Alexander D. M., Ward M., 2013, *ApJ*, 764, 159
 Liu Z. et al., 2016, *MNRAS*, 459, 1602
 López-Gonzaga N., Jaffe W., Burtscher L., Tristram K. R. W., Meisenheimer K., 2014, *A&A*, 565, A71
 López-Gonzaga N., Burtscher L., Tristram K. R. W., Meisenheimer K., Schartmann M., 2016, *A&A*, 591, A47
 Lusso E. et al., 2012, *MNRAS*, 425, 623
 Lusso E. et al., 2013, *ApJ*, 777, 86
 Maiolino R., Marconi A., Salvati M., Risaliti G., Severgnini P., Oliva E., La Franca F., Vanzani L., 2001, *A&A*, 365, 28
 Maiolino R. et al., 2010, *A&A*, 517, A47
 Marinucci A. et al., 2016, *MNRAS*, 456, L94
 Markowitz A. G., Krumpke M., Nikutta R., 2014, *MNRAS*, 439, 1403
 Mejía-Restrepo J. E., Lira P., Netzer H., Trakhtenbrot B., Capellupo D. M., 2018, *Nat. Astron.*, 2, 63
 Menzel M.-L. et al., 2016, *MNRAS*, 457, 110
 Merloni A. et al., 2014, *MNRAS*, 437, 3550
 Mor R., Netzer H., Elitzur M., 2009, *ApJ*, 705, 298
 Murray N., Quataert E., Thompson T. A., 2005, *ApJ*, 618, 569
 Pancoast A., Brewer B. J., Treu T., Park D., Barth A. J., Bentz M. C., Woo J.-H., 2014, *MNRAS*, 445, 3073
 Perna M., Lanzuisi G., Brusa M., Mignoli M., Cresci G., 2017, *A&A*, 603, A99
 Perola G. C. et al., 2004, *A&A*, 421, 491
 Peterson B. M. et al., 2004, *ApJ*, 613, 682
 Pol N., Wadadekar Y., 2017, *MNRAS*, 465, 95
 Ramos Almeida C. et al., 2009, *ApJ*, 702, 1127
 Ramos Almeida C. et al., 2011, *ApJ*, 731, 92
 Ricci C. et al., 2017, *Nature*, 549, 488
 Risaliti G., Elvis M., Nicastro F., 2002, *ApJ*, 571, 234
 Risaliti G., Salvati M., Marconi A., 2011, *MNRAS*, 411, 2223
 Schödel R. et al., 2007, *A&A*, 469, 125
 Scoville N., Norman C., 1995, *ApJ*, 451, 510
 Shen Y., Ho L. C., 2014, *Nature*, 513, 210
 Shen Y., Liu X., 2012, *ApJ*, 753, 125
 Shen Y., Greene J. E., Strauss M. A., Richards G. T., Schneider D. P., 2008, *ApJ*, 680, 169
 Shen Y. et al., 2011, *ApJS*, 194, 45
 Sulentic J. W., Zwitter T., Marziani P., Dultzin-Hacyan D., 2000, *ApJ*, 536, L5
 Treister E., Urry C. M., 2006, *ApJ*, 652, L79
 Tristram K. R. W., Burtscher L., Jaffe W., Meisenheimer K., Hönic S. F., Kishimoto M., Schartmann M., Weigelt G., 2014, *A&A*, 563, A82
 Vanden Berk D. E. et al., 2001, *AJ*, 122, 549
 Vasudevan R. V., Fabian A. C., 2007, *MNRAS*, 381, 1235
 Vasudevan R. V., Fabian A. C., 2009, *MNRAS*, 392, 1124
 Weingartner J. C., Murray N., 2002, *ApJ*, 580, 88
 Willott C. J., 2005, *ApJ*, 627, L101
 Wills B. J., Browne I. W. A., 1986, *ApJ*, 302, 56
 Zakamska N. L. et al., 2016, *MNRAS*, 459, 3144

APPENDIX A: HIGHLY OBSCURED SOURCES

Before the sample reselection, we exclude 11 highly obscured BLAGN2 with $N_{\text{H}} > 10^{23.5} \text{ cm}^{-2}$. Such sources have very different properties from the other BLAGN2, as shown in Fig. A1. They have significantly higher L_{X} but not accordingly higher M_{BH} . Conversely, most of them have much lower M_{BH} than the other BLAGN2. As a consequence, they appear as a high-end tail of the λ_{Edd} distribution. Their high λ_{Edd} might be a result of the effective Eddington limit, which increases with N_{H} (Fig. 3), in combination with the X-ray flux limit of the sample. However, in such highly obscured cases, the X-ray absorption correction on the basis of the *XMM-Newton* spectra (mostly below 10 keV) is highly model dependent. Their L_{X} are less reliable and can be overestimated. It is also possible that their M_{BH} are underestimated because of dust extinction of their optical emission.

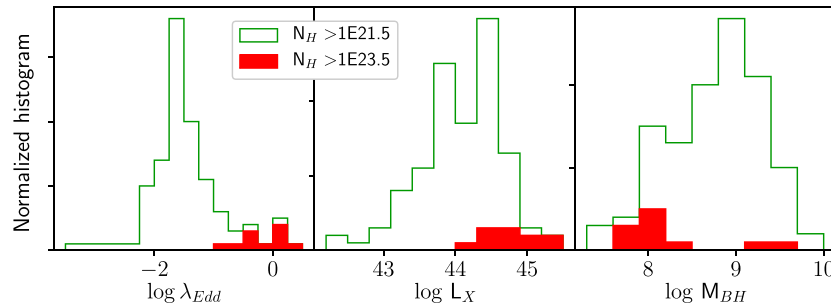


Figure A1. The λ_{Edd} , M_{BH} , and L_X distributions of the BLAGN2. The ones with $N_{\text{H}} > 10^{23.5} \text{ cm}^{-2}$ are filled with red colour.

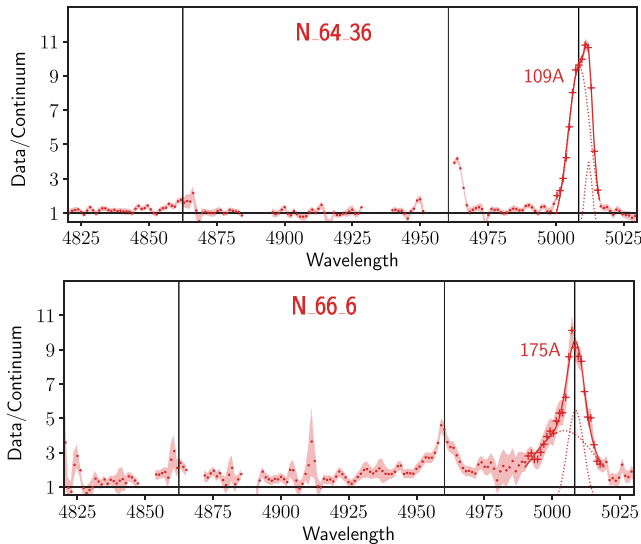


Figure B1. The ratio of data to local power-law continuum for the source N_64_36 and N_66_6. The same set of markers and lines is used as in Fig. 7. The EW of the [O III] 5007 line measured by the double-Gaussian fit is marked in the figure. Noise-dominated sections of the spectrum are eliminated.

APPENDIX B: OUTFLOWING SOURCES

As shown in Fig. 3, there are three BLAGN2 above the two-times-corrected effective Eddington limit of a low-dust-fraction absorber (the red dashed line, see Section 3.1 for details). Here we also consider the BLAGN2 that is below but the nearest to this limit. The ID (Liu et al. 2016) and redshifts of these four sources are N_89_36 at $z = 1.00$, N_66_6 at $z = 0.73$, N_64_36 at $z = 0.49$, and N_160_16 at $z = 2.34$, with λ_{Edd} from low to high. In such cases, unless the X-ray absorber is completely dust free or very far away from the black hole, it should be swept out by the radiation pressure (Fabian et al. 2009). In other words, outflow is expected. Therefore, we check whether their optical spectra show signs of outflow.

First, all of them have red ($\alpha' > 0$) optical continua (see Fig. 3), similar to other sources with outflow, which often show dust extinction (e.g. Brusa et al. 2015; Zakamska et al. 2016). Meanwhile, almost all the other sources in the dust blow-out region (blue solid line in Fig. 3) have $\alpha' < 0$.

Among the four sources, the two low-redshift ones could sample the $\text{H}\beta$ and [O III] wavelength range in their SDSS spectra. We plot

their spectra in this range in Fig. B1. We note that there are a lot of examples of outflowing sources found with no or very weak $\text{H}\beta$ line (e.g. Brusa et al. 2015; Kakkad et al. 2016). These two sources also show very weak $\text{H}\beta$ lines that are almost absent.

We fit the [O III] 5007 lines of the two sources with a double-Gaussian profile. As shown in Fig. B1, both their [O III] 5007 lines present asymmetric shapes with strong outflowing (blueshifted) components,⁴ similar to what is conventionally used to select objects with outflows (e.g. Harrison et al. 2014; Perna et al. 2017). We argue that, in such cases, it might be the outflowing polar dust that reddens the optical continua and weakens the broad $\text{H}\beta$ line.

APPENDIX C: COMPOSITE SPECTRUM OF BLAGN

Although, theoretically, geometric mean spectrum has an advantage of preserving the power-law continuum shape over median spectrum, this advantage is impractical in practice, because the spectra are not always power laws – a small fraction of them show a reddening caused by dust extinction (see Fig. 2). The median composite spectrum is of more interest and has been used as a cross-correlation template. In Fig. C1, we show the median composite spectrum of our BLAGN in comparison with the median composite quasar spectrum obtained by Vanden Berk et al. (2001).

Our composite spectrum shows stronger emission lines, a flatter power law below 4000 \AA , and a red excess above 4000 \AA . All these differences are caused by different sample selections – Vanden Berk et al. (2001) used colour-selected quasars but our BLAGN are selected on the basis of X-ray brightness and optical emission lines. First, our sample tends to select sources with stronger emission lines. Meanwhile, we could select the BLAGN in spite of moderate dust extinction of the continuum emission from the disc. Such dust-reddened sources are responsible for the flatter power law of our composite spectra. The red excess above 4000 \AA corresponds to a stronger stellar component. The relative strength of the stellar component is various among different samples because it depends on fibre diameter, redshift, and AGN luminosity (see also the discussion in Pol & Wadadekar 2017). Excluding the significantly reddened BLAGN with $\alpha' > 0$, the composite spectrum of our BLAGN (blue line and points in Fig. C1) is more similar to that of Vanden Berk et al. (2001).

⁴The asymmetric shape of [O III] is causing a slight underestimation of redshift of N_64_36.

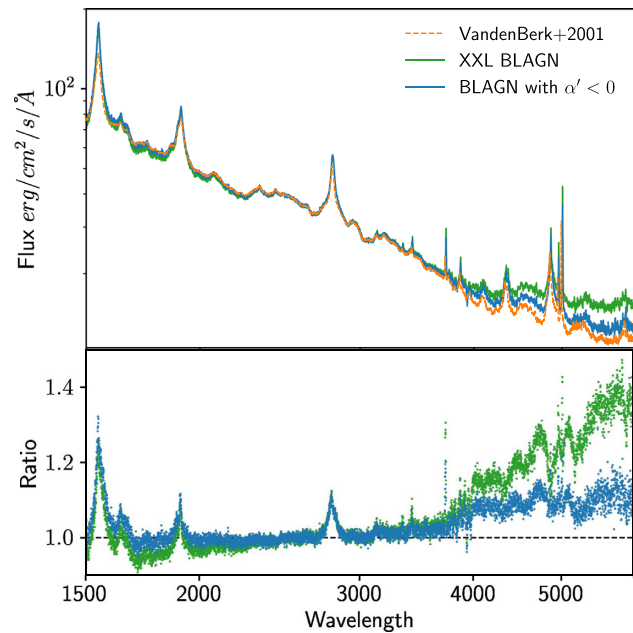


Figure C1. The median composite spectrum generated using method ‘A’ (Section 4.2) for our BLAGN (sample ‘1’, green) and for the blue ($\alpha' < 0$) subsample (blue). The 68 per cent confidence intervals are very narrow and thus not shown. For comparison, we plot the composite quasar spectrum of Vanden Berk et al. (2001) (orange dashed line). The spectra are normalized between 3020 and 3100 Å. The lower panel shows the ratios of our composite spectra to that of Vanden Berk et al. (2001).

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