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Insoluble Organic Matter in Chondrites: Archetypal Melanin-Like PAH-Based Multifunctionality at the Origin of Life?

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Abstract

An interdisciplinary review of the chemical literature that points to a unifying scenario for the origin of life, referred to as the Primordial Multifunctional organic Entity (PriME) scenario, is provided herein. In the PriME scenario it is suggested that the Insoluble Organic Matter (IOM) in carbonaceous chondrites, as well as interplanetary dust particles from meteorites and comets may have played an important role in the three most critical processes involved in the origin of life, namely *1) metabolism*, via *a)* the provision and accumulation of molecules that are the building blocks of life, *b)* catalysis

(e.g., by templation), and c) protection of developing life molecules against radiation by excited state deactivation; 2) **compartmentalization**, via adsorption of compounds on the exposed organic surfaces in fractured meteorites, and 3) **replication**, via deaggregation, desorption and related physical phenomena. This scenario is based on the hitherto overlooked structural and physicochemical similarities between the IOM and the dark, insoluble, multifunctional melanin polymers found in bacteria and fungi and associated with the ability of these microorganisms to survive extreme conditions, including ionizing radiation. The underlying conceptual link between these two materials is strengthened by the fact that primary precursors of bacterial and fungal melanins (collectively referred to herein as *allomelanins*) are hydroxylated aromatic compounds like homogentisic acid and 1,8-dihydroxynaphthalene, and that similar hydroxylated aromatic compounds, including hydroxynaphthalenes, figure prominently among possible components of the organic materials on dust grains and ices in the interstellar matter, and may be involved in the formation of IOM in meteorites. Inspired by this rationale, a vis-à-vis review of the properties of IOM from various chondrites and non-nitrogenous allomelanin pigments from bacteria and fungi is provided herein. The unrecognized similarities between these materials may pave the way for a novel scenario at the origin of life, in which IOM-related complex organic polymers delivered to the early Earth are proposed to serve as PriME and were preserved and transformed in those primitive forms of life that shared the ability to synthesize melanin polymers playing an important role in the critical processes underlying the establishment of terrestrial eukaryotes.

Keywords: insoluble organic matter; melanins; polycyclic aromatic hydrocarbons (PAHs); Primordial Multifunctional organic Entity (PRiME)

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1. Introduction

Meteorites, with the exception of a few samples from Mars and the Moon, are fragments of asteroids from the asteroid belt located between Mars and Jupiter. The chondritic meteorites are the most primitive type of meteorite having avoided the melting and differentiation that affected other meteorite types. The heating of planetesimals in the early Solar System was mostly driven by the decay of the short-lived radionuclides ^{26}Al ($t_{1/2} \approx 0.7$ Ma) [1]. The chondrite parent bodies formed between ~ 2 Ma and ~ 4 Ma after Solar System formation when there was still enough ^{26}Al to drive rock-forming geological processes (lithification), such as thermal metamorphism and aqueous alteration, but not enough to melt them. Based on their bulk compositions and physical properties, the chondrites have been subdivided into a number of classes and groups: enstatite (EH and EL), Rumuruti, ordinary (H, L and LL), and carbonaceous (CI, CM, CO, CV, CK, CR, CH and CB) [2-5]. Chondrites are also assigned petrologic types (e.g., CM1 and LL6) depending on the extent of aqueous alteration (1-3, with 1 being the most aqueously altered, and 3 the least) or thermal metamorphism (3-6, with 6 being the one that experienced the most thermal metamorphism) they experienced. Carbonaceous chondrites are a pristine class of materials that formed from the protoplanetary disks and together with interplanetary dust particles (IDPs) from meteorites and comets and Ultra-Carbonaceous Antarctic Micrometeorites (UCAMMs) provided the major source of organic carbon in the early Earth.

Chondrites contain inclusions (chondrules and refractory inclusions) set in a fine-grained matrix. It is in the matrix that one finds the most primitive materials, such as presolar circumstellar grains and organic matter. The circumstellar grains are identified based on their large isotope anomalies that can only be explained by nucleosynthetic processes. Ices also seem to have been important components of the matrices of all chondrites, except the highly reduced enstatite chondrites. At the time of the formation of their parent bodies, the matrices of all chondrites may have contained ~ 3 -4 wt.% of organic carbon. This organic material is generally divided in a complex suite of solvent Soluble Organic Matter (SOM) and a more abundant macromolecular component, the Insoluble Organic Matter (IOM). IOM is a

descriptive term that has been widely adopted to indicate the larger portion (75–95%) of the organic carbon present in chondrites [6,7]. The remaining % of the organic matter in carbonaceous chondrites is composed of a mixture of solvent-soluble organic compounds (i.e., SOM) [8-10]. The organic material in most chondrites has been heavily modified or destroyed by thermal metamorphism and shock heating. However, members of the CI, CM and CR groups experienced variable degrees of low temperature (<100-150°C) aqueous alteration. The extent to which aqueous alteration has modified the SOM and IOM remains controversial. Nevertheless, the IOM and many soluble organic compounds in these meteorites have large excesses in deuterium, ^{13}C and ^{15}N [11-14]. D and ^{15}N enrichments suggest that they or their precursors formed in low temperature (10-40 K) and radiation-rich environments in the interstellar medium (ISM) or the outer Solar System [15-23]. Also, CI- and/or CM-like materials are the leading candidates for the major sources of volatiles (e.g., hydrogen, carbon and nitrogen) in the early Earth. If some of this material was delivered by chondritic IDPs, UCAMMs, micrometeorites and meteorites, then they would have also been a significant source of complex, prebiotic organic matter.

Due to its insolubility, IOM has been difficult to characterize in detail at the molecular level and is known mostly through spectroscopic, physical and chemical/pyrolytic analyses. IOM-like materials are also present in various geochemical contexts, such as sedimentary rocks, and are referred to as kerogens.

Despite increasing interest in chondrites as rich sources of soluble and volatile organic compounds of possible relevance to the origin of life, relatively little attention has been devoted to its much more abundant insoluble counterpart. Apart from some systematic chemical and structural characterization studies [6], the general attitude in the literature on the origin of life has been to disregard IOM as an essentially inert material and favor the richer and more readily appreciable chemistry of soluble organic compounds.

The aim of this paper is to draw the attention of scientists involved in research on astrobiology and the origin of life to the so far overlooked, yet most relevant observation that *most if not all of the key*

structural features and properties of the IOM found in carbonaceous chondrites are shared by melanins, the ubiquitous dark insoluble phenolic polymers found in almost all types of living organisms, including lower level organisms like bacteria and fungi. Such an analogy emerges clearly from a background of scattered observations, a virtually unprecedented integration of astrochemical, astrophysical, organic and physical chemical viewpoints, and a critical perspective of the evolutionary significance of melanin and related complex organic biopolymers. From bacteria to fungi, plants, cephalopods, birds, mammals up to humans, melanins appear to serve a broad variety of roles and functions which can be re-conceptualized and revisited as *fulfilling most of the key chemical requirements at the origin of life*. Analogies and matching properties between IOM and melanins in bacteria and fungi, both natural and synthetic, are used herein to suggest that IOM may have been *an active player in the origin and processing of life molecules* and thus may provide the basis to complement and integrate current prebiotic scenarios in the early Earth into a unifying perspective of possible evolutionary relevance.

2. Melanin-type materials and precursors in astrochemical environments

The plausibility that an archetypal melanin-like material could have played a significant role in the emergence of molecular complexity is also based on the availability of suitable chemical precursors in astrochemical environments. The chemical precursors of melanins in extant taxa are represented by two class of organic compounds: *i*) the aromatic amino acids phenylalanine (Phe) and tyrosine (Tyr) along with the sulfur containing derivative cysteine (Cys) (Fig. 1, panel A), that are involved in the synthesis of nitrogen containing mammalian melanins, eumelanins and pheomelanins, and *ii*) the aromatic naphthols 2,3,6,8-tetrahydroxynaphthalene, 2,4,8-trihydroxynaphtalene and 1,8-dihydroxynaphtalene (Fig. 1, panel B), involved in the synthesis of fungal allomelanins. Tyrosine is also involved in the synthesis of the bacterial allomelanins, pyomelanins, via the intermediate formation of homogentisic acid.

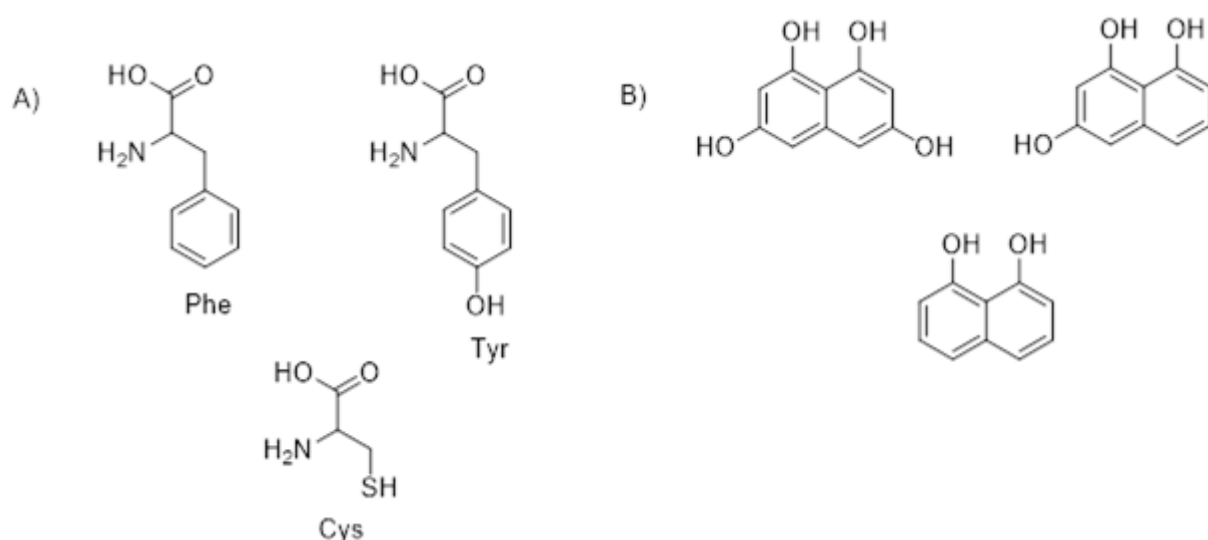


Fig. 1. Chemical precursors of melanins. Panel A: phenylalanine (Phe), tyrosine (Tyr) and cysteine (Cys) amino acids as chemical precursors of eumelanins and pheomelanins. Panel B: 2,3,6,8-tetrahydroxynaphthalene, 2,4,8-trihydroxynaphthalene and 1,8-dihydroxynaphthalene as chemical precursors of fungal allomelanins.

Besides biological environments, melanin-precursor-forming reactions may also operate in a wide range of space environments that are subject to different kinds of conditions, including irradiation by stellar winds, cosmic rays, as well as thermal heating by radioactive decay [24] and aqueous alteration. Then the question is: are the chemical precursors of melanin abundant enough in space to be compatible with the presence of melanin-type materials in astrochemical environments?

Chemical precursors of melanins have been detected in the SOM of carbonaceous chondrites. The SOM of carbonaceous meteorites has a high molecular diversity (in the order of tens of thousands), with a bulk composition of $\sim\text{C}_{100}\text{H}_{155}\text{O}_{20}\text{N}_3\text{S}_3$, [25]. It includes amino acids, diamino acids, dipeptides, hydrocarbons, alcohols, carboxylic acids, diketopiperazines, sulfonic and phosphonic acids, purines, pyrimidines, sugars and sugar-related compounds, amines, amides, aldehydes, and ketones [8]. The extraterrestrial soluble organic molecules present in carbonaceous chondrites are enriched in ^{13}C , D and ^{15}N [14, 26-31], which indicates an interstellar heritage of these molecules or, most likely of their precursors [15-21].

Amino acids have been extensively analyzed in carbonaceous chondrites since the 1970's, with more than 80 different amino acids identified (carbon number from C₂ through C₉), and total amino acid abundance of around 60ppm. These have a complete structural diversity, and a decrease in concentration with increasing carbon number [32-36]. While several carbonaceous chondrites were analyzed for amino acids, including CMs, CIs, CRs, CVs, COs, and CKs [37-67], the aromatic amino acids Phe and Tyr are only present in CR chondrites, which are reported in Table 1 [68,69]. To note however, that this last study does not include any procedural blanks or measurement errors, which are crucial to address the accuracy of the measurements [70]. To date, no compound specific isotopic compositions have been performed for Phe or Tyr, and therefore it is not possible to determine without a question whether these two amino acids are indigenous to the carbonaceous chondrites or terrestrial contamination [8].

Table 1. Abundance of phenylalanine and tyrosine in the hydrolyzed, water-extracts of different CR chondrites (nmole/g of meteorite) [104, 105].

Meteorite^a	D,L-Phenylalanine	D,L-Tyrosine
MET 00426	17.8	8.0
GRO 95577	≤1	≤1
MIL 07525	nf ^b	12.8
PCA 91082	nf ^b	≤1
QUE 99177	nf ^b	≤1
EET 92042	8.5	3.0
GRA 95229	24	2
LAP 02342	127	≤1

^aThe Renazzo-type (CR) carbonaceous chondrites: Meteorite Hills (MET) 00426, Grosvenor Mountains (GRO) 95577, Miller Range (MIL) 07525, Pecora Escarpment (PCA) 91082, Queen Alexandra Range (QUE) 99177, Elephant Moraine (EET) 92042, Graves Nunataks (GRA) 95229, and LaPaz Icefield (LAP) 02342. ^bNot found.

Sulphur containing amino acids have not been observed in carbonaceous chondrites; however, products of sulfhydryl oxidation, such as sulfonic acids and alkyl sulfonates, have been detected in the Murchison meteorite [71], indicating the occurrence of an extensive chemical processing [72].

Seven of the eight possible alkyl sulfonic acids were identified in this CM2 chondrite, with carbon atoms up to C₄, abundances decreasing with increasing carbon number (highest concentration from around 380 nmol/gram) [71]. ¹³C and D chemical shift values were obtained for alkyl sulfonic acids present in the Murchison meteorites: ¹³C chemical shift values of +29.8‰, +9.1‰, -0.4‰, and -0.9‰ for the methyl, ethyl, isopropyl, and n-propyl sulfonic acids, respectively, and D chemical shift values of +483‰, +787‰, +536‰, and +852‰ for the methyl, ethyl, isopropyl, and n-propyl sulfonic acids, respectively [14].

Since the pivotal Miller experiment [73], different prebiotic processes under both space- and planetary-like conditions have been described for the abiotic synthesis of amino acids [74], including Phe, Tyr and Cys derivatives [75]. The abiotic synthesis of Phe and Tyr occurs in a large variety of space and geochemical scenarios, including meteorite-catalyzed thermal processes [76], spark discharge in gas mixtures [77], associated with self-crystallization and racemic resolution [78], high pressure hydrothermal transformations [79], and the geological process of serpentinization [80], also focusing on the catalytic role of biomorph structures [81]. The possible prebiotic synthesis mechanisms for Cys include irradiation [82] and hydrothermal processes [83]. The latter mechanism is generally thought to involve so-called Strecker-cyanohydrin synthesis, a sequential chemical condensation of hydrogen cyanide (HCN) and aldehydes [84]. Recently, the major role of HCN in the formation of amino acids under spark discharge conditions has been revisited, and formamide (NH₂CHO), a widely distributed one-carbon compound in the universe [85], has become the favored intermediate in the process [86]. Examples of the catalytic effects of meteoritic minerals in the prebiotic synthesis of amino acids (and of other biologically relevant molecules) from formamide have been reported for a large variety of physical and chemical conditions. The synthesis could have been fueled by a range of energy sources (thermal,

solar wind, cosmic rays, radiolysis) [76] when compatible with the stability of the end products [87]. Reductive amination of α -keto acids and carbonylation of amines under proton irradiation [88] and Fisher-Tropsch-type (FTT) syntheses [89] have also been proposed as alternative synthetic pathways, both inside and outside (the icy surfaces) [89] the meteorite parent bodies. Although the detailed mechanism(s) for the formation of aromatic amino acids needs to be further explored, a Friedel-Craft type functionalization of a pre-formed benzene ring with acetylene followed by electrophilic addition reactions and hydrolysis has been suggested as possible reaction pathway for the synthesis of Phe and Tyr (Fig. 2) [90].

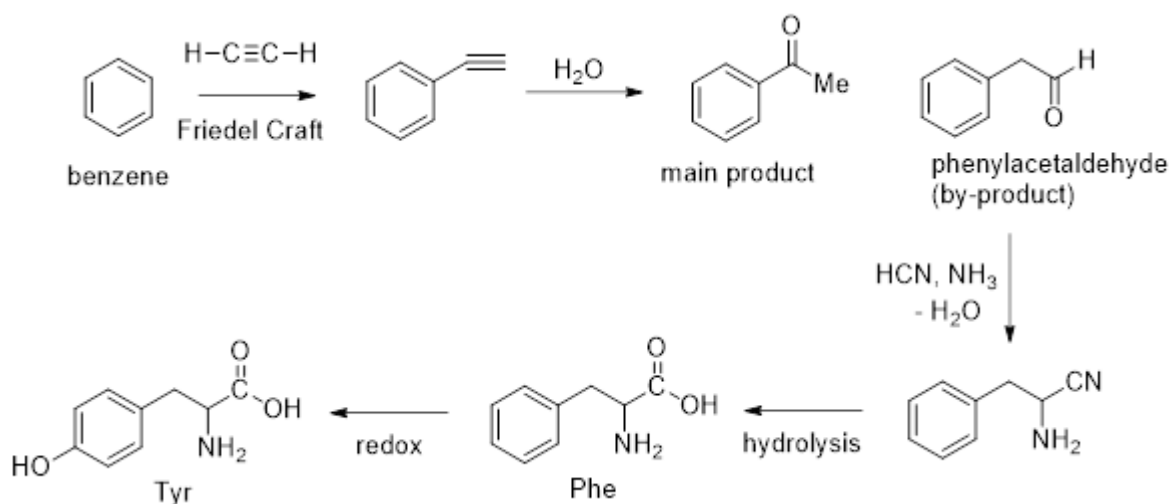


Fig. 2. Abiotic synthesis of phenylalanine (Phe) and tyrosine (Tyr) as suggested by Friedman and Miller [92].

As mentioned previously, allomelanins are mainly produced by oxidative polymerization of DHN, a dihydroxylated derivative of the simplest PAH naphthalene (C_{10}H_8) [91]. A large number of PAHs have been detected in carbonaceous chondrites, including naphthalene and alkyl-naphthalene derivatives [55,92-96]. The astrochemical interest in PAHs derives from their detection in various astronomical environments, such as cosmic dust, icy satellites, carbonaceous chondrites, comets. Furthermore, their possible role in prebiotic processes led to introduction of the so-called “PAH world hypothesis” [97].

While aromatic naphthols have not been detected in carbonaceous chondrites, naphthalene has been detected in Y-791198 [93], Y-74662 [94] and Murchison [95,96]. Structural isomers of methyl- and dimethylnaphthalene have been detected in Y-791198 [93], and Murchison [95]. The abundances of alkylated PAHs in CM2 chondrites seem to be related to the extension of aqueous alteration on its parent body [55,98]. For example, the Paris meteorite, one of the least aqueously altered carbonaceous chondrite analyzed to date does not contain any alkylated PAHs [55], while other more aqueously altered CM2s (like Murchison and Y-791198) contain alkylated naphthalene.

The Rosetta Orbiter Spectrometer for Ion and Neutral Analysis (ROSINA) analyzed comet 67P/Churyumov-Gerasimenko and detected toluene, but no naphthalene or alkylated naphthalene [99]. The smallest amino acid glycine was also detected in this comet [100], as well as in Wild-2 comet [101,102], with no larger amino acids indisputably being detected. The SOM content of Antarctic micrometeorites (AMMs) has also been analyzed, in particular for PAHs [103,104], and amino acids [105-107]. The AMMs were strongly depleted in naphthalene relative to Murchison [106], and had very low abundances of amino acids [105-107].

3. Insoluble Organic Matter from chondrites and comets

3.1. Definition and occurrence

The IOM is the major carbonaceous component in chondrites that has been isolated and characterized [7,10]. Two basic techniques have been used to demineralize meteorites and concentrate the IOM, one utilizing HF-HCl and the other using CsF-dioxane. Both techniques produce similar IOM yields and compositions. However, the IOM recovered in demineralized residues typically only accounts for ~60% of the bulk carbon [108]. Only a small fraction of the unrecovered carbon can be accounted for by carbonate or solvent extractable organic material. The remainder is probably either in very fine-grained IOM that is hard to recover from the demineralizing solutions, or is in acid hydrolysable functional groups that are attached to the IOM. However, at present we cannot exclude that some of

the missing carbon is present in soluble compounds that are strongly associated with minerals and as a result they cannot be efficiently extracted with solvents prior to demineralization.

In situ observations of carbonaceous grains in matrix, which are presumably dominated by IOM, are free of mineral inclusions and exhibit no obvious spatial relationships with any minerals [109]. The grains have a diversity of morphologies, sizes (mostly $< 1\mu\text{m}$) and isotopic compositions [109,110]. These include isolated irregular grains and roughly spherical and often hollow nanoglobules, as well as veins. The veins probably reflect a redistribution of material when ices in the matrix melted and the fluids began to move [111].

This diversity of grain morphologies and isotopic compositions is preserved in demineralized IOM isolates, and it is these residues that have been studied in the most detail. IOM in the CR chondrites appears to be particularly primitive and its bulk elemental composition, $\text{C}_{100}\text{H}_{75-79}\text{O}_{11-17}\text{N}_{3-4}\text{S}_{1-3}$ [10], closely resembles the bulk elemental compositions of comet Halley CHON particles, $\sim\text{C}_{100}\text{H}_{80}\text{O}_{20}\text{N}_4\text{S}_2$ [112], and the refractory organics in comet 67P/Churyumov-Gerasimenko dust particles, $\sim\text{C}_{100}\text{H}_{100}\text{N}_3$, measured by the COSIMA instrument on the Rosetta spacecraft [113,114]. The refractory organics makes up $\sim 30\text{-}40$ wt.% of these two comets [115-117], roughly an order of magnitude higher than in chondrite matrices. The CR2s appear to have experienced the most benign alteration conditions for IOM of any chondrite group. The IOM from all other chondrite groups tend to have lower H/C ratios that roughly reflect the intensities of the parent body processing that their host meteorites experienced [10]. However, the extent to which the variations in IOM elemental and isotopic compositions between the most primitive members of each chondrite group were inherited or are the products of parent body processes remains debated [118-120]. The bulk IOM in CR chondrites is also, with rare exceptions, the most isotopically anomalous, both in bulk [10] and at micron to submicron scales [118,121,122]. When analyzed at similar spatial scales, the range of hydrogen and nitrogen isotopic compositions seen in the most primitive IOM is very similar to that seen in the most primitive interplanetary dust particles collected in the stratosphere that may be of cometary origin [123,124].

3.2. Structural and physical properties

Chemical degradation [125-128] and pyrolysis [129-132] studies point to the IOM being composed of small (1-4 ring) polyaromatic units (PAUs) that are highly substituted and cross-linked by short, highly branched aliphatic and ester linkers (Fig. 3). However, yields from both chemical degradation and pyrolysis studies are low, so it is not clear how representative these chemically and thermally accessible components are of the bulk material. In addition, the yields of individual compounds are rarely determined quantitatively. IR spectroscopy suggests that at least the aliphatic material in bulk IOM is indeed composed of short and highly branched carbon units. Based on RuO_4 oxidation experiments, it was concluded that the aliphatic material forms bridges between aromatic units [130]. The high degree of cross-linking in IOM likely explains its refractory nature.

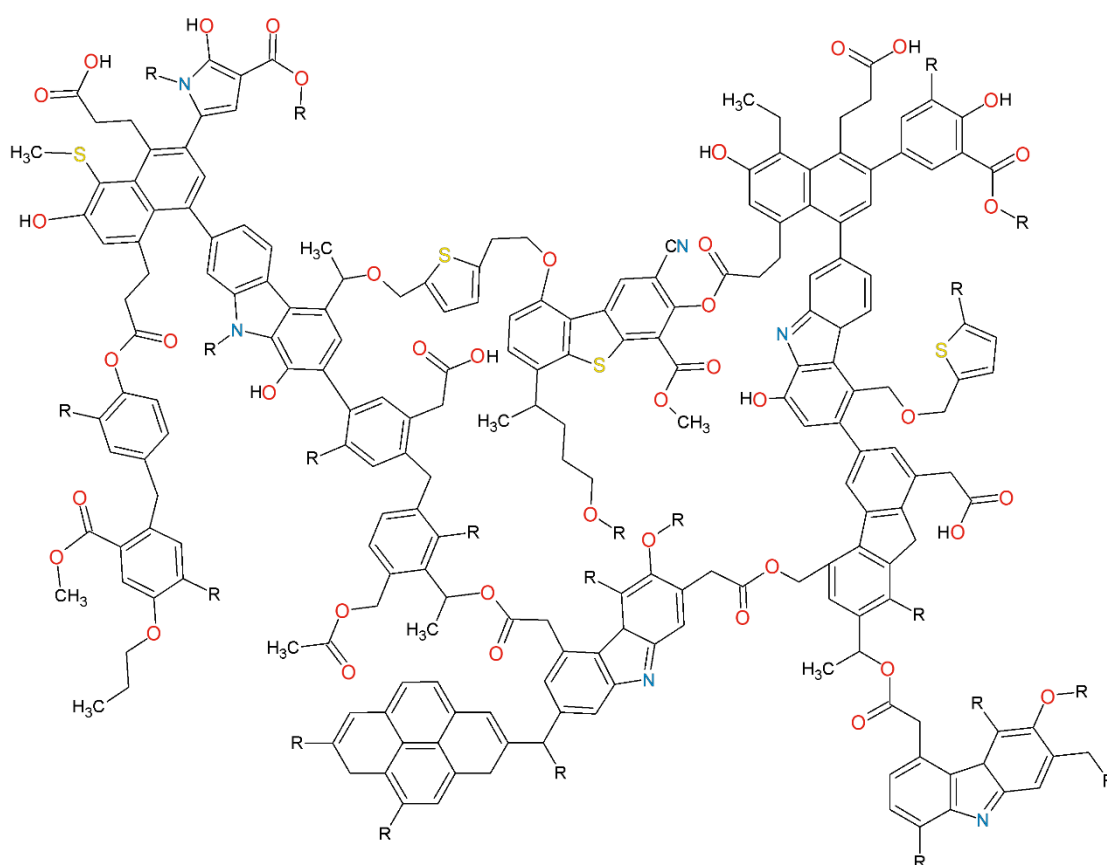


Fig. 3. Molecular model for the chemical structure of Murchison IOM. *R* denotes an organic moiety (Modified from Derenne and Robert (2010) [133,134]). This model cannot be considered as a detailed representation of the molecular structure of the IOM, but it rather illustrates the main characteristics of the C skeleton and the O-, N- and S-containing groups.

Because of the low yields and poor quantification, at present Nuclear Magnetic Resonance (NMR) provides the best quantitative picture of the main functional group in the bulk IOM [135-140]. ^1H - and ^{13}C -NMR also suggest that most of the PAUs in IOM is small and highly substituted, and that the aliphatic component be composed of short, highly branched units. There is considerable line broadening in the NMR spectra. While some of this could be due the presence of paramagnetic minerals like chromite [141], most of it certainly reflects the considerable chemical complexity. Such complexity is perhaps not too surprising given that the analyzed IOM samples are collections of very large numbers of morphologically and chemically heterogeneous grains. In addition, the NMR measurements indicate that the oxygen functionality is dominated by substituted furan/pyran moieties, highly branched oxygenated aliphatics and carbonyl groups. X-ray Absorption Near Edge Structure (XANES) spectroscopy [109,110, 142] and electron energy loss spectroscopy (EELS) [143] applied at the carbon K-edge also provides a valuable tool for assessing the functional group chemistry of IOM at scales that range down to 100 nm or less. C-XANES spectra exhibit a strong absorption feature corresponding to $\text{C}=\text{C}$ groups, likely related to aromatic rings, along with absorption corresponding to carboxylic groups. Minor absorption features due to ketone/phenol/nitrile and hydroxyl groups are also seen. As aliphatic carbon is a weak absorber at the carbon K-edge, its contribution to the spectra is faint, but noticeable.

To date, a ^{15}N -NMR spectrum has only been reported for Orgueil IOM [130]. It shows a strong contribution for pyrrole groups, along with minor contribution from nitrile. However, studies of IOM from other primitive meteorites have failed to see any ^{15}N -NMR signal above background despite having similar nitrogen contents to Orgueil IOM [130,144]. Nitrogen functionality can also be probed by XANES spectroscopy. The nitrogen K-edge XANES spectra are remarkably similar across chondrite groups and petrologic types, suggesting that the nitrogen functional groups are very stable. The spectra are also remarkably featureless, with only a few very small peaks despite the transitions for many nitrogen functional groups being very intense. One peak is consistent with a very small

amount of imine and/or nitrile. Another peak is consistent with pyrrolic, amine or amide nitrogen. Otherwise, the near continuous absorption spanning the entire N-XANES region requires the presence of a range of as yet unidentified nitrogen functional groups.

XANES also provides the best evidence for the sulfur speciation in IOM, although it has been the subject of only one study of bulk IOM [145]. In that study, the sulfur functionality was divided into three major groups with similar absorption energies: aliphatic sulfur (disulfides, thioethers and thiols), sulfur heterocycles (thiophenes and thianthrenes) and oxidized sulfur groups (sulfoxides, sulfones and sulfonates). The most intense absorption features in all the IOM spectra are due to aliphatic sulfur, with much weaker heterocycle and oxidized sulfur intensities. The relative heterocycle and oxidized sulfur intensities are higher in more aqueously altered meteorites, suggesting that there was some parent body modification of the sulfur functionality.

The Tagish Lake meteorite has a number of distinct lithologies that experienced varying extents of hydrothermal alteration, probably associated with a brief heating event(s). The IOM in these lithologies exhibit varying degrees of aromaticity (and H/C) that correlate with the extents of alteration, suggesting that the two motifs are linked [146,147]. The varying degrees of aromaticity seem to have been produced with little or no loss of carbon or nitrogen, although there was loss of hydrogen and isotopic evolution due to interaction with the aqueous fluid. The apparently facile and efficient conversion of aliphatic to aromatic is, perhaps, best explained if the aliphatic material in the original IOM was already organized in cyclic structures able to minimize the number of bond breaking and restructuring. This is consistent with observations in various CM chondrites of increasing alteration degrees; the IOM in more altered CMs shows more aromatic carbon, less aliphatic chains [135,136] and a decreasing H/C ratio.

The hydrogen isotopic compositions of individual pyrolysates (both aromatic and aliphatic) and aliphatic side groups on the PAUs released by RuO₄ are comparable to those of the bulk IOM compositions, albeit with significant scatter [128,148,149]. It has been suggested that this scatter is the result of variations in the average hydrogen isotopic compositions of moieties in the IOM, reflecting

heterogeneities at the molecular scale that could be correlated with H-C bond energies [149]. This correlation has been reproduced by irradiation experiments on simple organic polymers [150]. This interpretation is not consistent with all pyrolysis results [148]. On the other hand, electron paramagnetic resonance (EPR) experiments suggest that radicals in IOM are associated with very extremely D-rich benzylic positions on aliphatic side chains [151,152] that could eventually constitute localized micron-sized D-rich enrichments called hotspots [23]. The C-H bond energies of benzylic radicals is amongst the lowest in the IOM. The EPR experiments also show that the radical speciation in IOM (dominated by diradicals/diradicaloids) is quite unlike those of terrestrial coals and kerogens (exclusively monoradicals) despite their superficial resemblance in terms of elemental and functional group chemistries [153]. Irradiation by energetic particles in the ISM or solar protoplanetary disk is one possible mechanism for generating these radicals. Irradiation experiments of simple polymers have successfully produced monoradicals and diradicals but at much lower concentrations than observed in IOM [154]. Unfortunately, that study was unable to determine if diradicaloids were present after the irradiation, leaving their origin unresolved.

3.3. Origin

Currently, there is no consensus about where, how and from what the IOM formed [112,7]. One possibility that was originally put forward because of the large deuterium and ^{15}N enrichments in IOM is that it or its precursors formed in the protosolar molecular cloud [9,10,16,155,156]. It has long been known that simple molecules synthesized in the relatively dense (10^2 - 10^6 particles/cm³), cold (10-30 K), radiation-rich environments of molecular clouds can become enormously enriched in deuterium relative the dominant H₂ [157]. The circumstellar grains found in primitive meteorites clearly survived Solar System formation. The lifetimes of circumstellar grains in the ISM is relatively short, so an even larger amount of dust that formed in the ISM is also likely to have survived Solar System formation. The ISM silicate dust should have an average isotopic composition that is like that of the bulk Solar System, making it hard to identify. This may not be the case for carbonaceous dust. Refractory

carbonaceous dust and large PAHs make up a significant fraction of the dust in the diffuse ISM [158,159]. While there are striking similarities in the aliphatic C-H stretch regions of the IR spectra of the diffuse ISM dust and IOM [160,161], the diffuse ISM dust does not seem to contain significant amounts of heteroatoms (O, N and S), unlike IOM. What happens to the carbonaceous dust once it is incorporated into the cold (10-30 K), dense regions of molecular cloud where star-formation occurs is still a matter of debate. One possibility is that the grains would become coated in complex interstellar ices (bearing H₂O, CO, CO₂, H₃COH, H₂CO, HCN, NH₃, NH₂COH etc.) and irradiated by energetic cosmic rays [162,163]. Radiation damage by the cosmic rays would generate defects and radicals in the grains and ices. When the grains/ices are warmed during star formation, these radicals become mobile and react adding heteroatoms to the carbonaceous grains. The H₂O and many of the other compounds in the ices are known to be deuterium- and ¹⁵N-rich, so this mechanism would naturally explain the isotopic composition of the IOM.

An origin in the ISM would imply that a common IOM precursor was incorporated into cometary and asteroidal parent bodies, and that in the meteorite parent bodies, at least, it was subsequently altered under a range of conditions [10,164]. Whether all the elemental and isotopic variations found amongst the chondrites can be attributed to parent body processes remains debated. It is possible that the IOM precursor material was modified to varying degrees in the solar nebula prior the accretion of the chondrite parent bodies [118,165], but at present clear evidence that this happened is lacking.

Alternatively, the IOM or its precursors could have formed in the solar nebula. Formation of IOM by FTT synthesis was the first widely accepted mechanisms for its origin [166,167]. However, there is no spatial association between IOM and plausible mineral catalysts (Fe,Ni-metal, magnetite, etc.).

Moreover, FTT synthesis is likely to have been efficient only in a relatively narrow region of the inner Solar System [167], which would make it difficult to account for the high abundances of refractory organics in comets. In addition, experiments suggest that under nebular conditions the dominant products of FTT synthesis would be volatile gases [168,169]. The isotopic composition of the IOM also cannot be explained by FTT synthesis alone.

A nebular alternative to FTT synthesis is irradiation of ices by cosmic rays and stellar radiation when they are lofted to high altitudes in the disk [170]. To date, ice irradiation experiments have not produced IOM-like material. It is possible that irradiation of organic residues after sublimation of the ice can accomplish this, but experimental evidence that this is the case is lacking. Such a nebular irradiation scenario may be rather similar to how dust is cycled between the diffuse and dense ISM, but in the solar nebula it may be harder to generate the large deuterium and ^{15}N anomalies found in IOM. It has been suggested that the deuterium enrichments, at least, could be the result of interaction of carbonaceous grains with cosmic ray generated H_2D^+ [149,151]. In molecular clouds, H_2D^+ is one of the species responsible for producing highly deuterated organic molecules, but it is not clear that cosmic ray fluxes would have been sufficient to generate much H_2D^+ in the solar nebula because strong stellar winds would have shielded the disk [171,172].

One mechanism that has arguably been explored experimentally more than any other is polymerization of interstellar formaldehyde from solution via the formose reaction when ices melt in meteorite or cometary parent bodies [140,173,174]. To reduce reaction times, most experiments have included glycolaldehyde as a starting product as it is an early intermediate in the formose reaction. The process naturally produces precipitates that in size and morphology look very similar to the nanoglobules found in chondrites and IDPs. At low temperature, the product is much more oxygen-rich than IOM, but heating it briefly to $\sim 250^\circ\text{C}$ can generate a material whose NMR spectrum resembles in many ways IOM. Despite these promising results, this mechanism must overcome significant objections. The formose reaction is relatively inefficient with yields of at best $\sim 20\%$ and formaldehyde is a minor component of ISM and cometary ices [59]. Similar experiments conducted on hexamethylenetetramine (HMT), recently identified in the carbonaceous chondrites Murchison, Murray and Tagish Lake [175], result in the formation of a N-rich insoluble organic material but with an even smaller yield (1%) [176]. This makes the abundances of IOM in chondrites and the related refractory organics in comets problematic. It is also unclear that the volatile-rich (CO , CO_2 , etc.) ices of comets like 67P/Churyumov-Gerasimenko could ever have been molten.

3.4 Formation pathways of putative IOM precursors and aromatic compounds

Illustrated schematically in Fig. 4 is the standard picture of how small carbon-bearing molecules, such as acetylene, ethylene and methane, that are formed in the outflows of carbon stars can give rise to PAHs and then to carbonaceous particles through combustion-like processes [177,178]. However, a robust explanation of how these large molecules form is still lacking, especially in evolved stars where the characteristic aromatic infrared bands are only detected at late stages. Indeed, recent experimental work found that amorphous carbon nanograins and aliphatic carbon clusters can form through low-pressure gas-phase condensation of carbon atoms in a hydrogen atmosphere with a C/H₂ ratio that is similar to that reported for evolved stars, while aromatic species or fullerenes do not form effectively under these conditions [179].

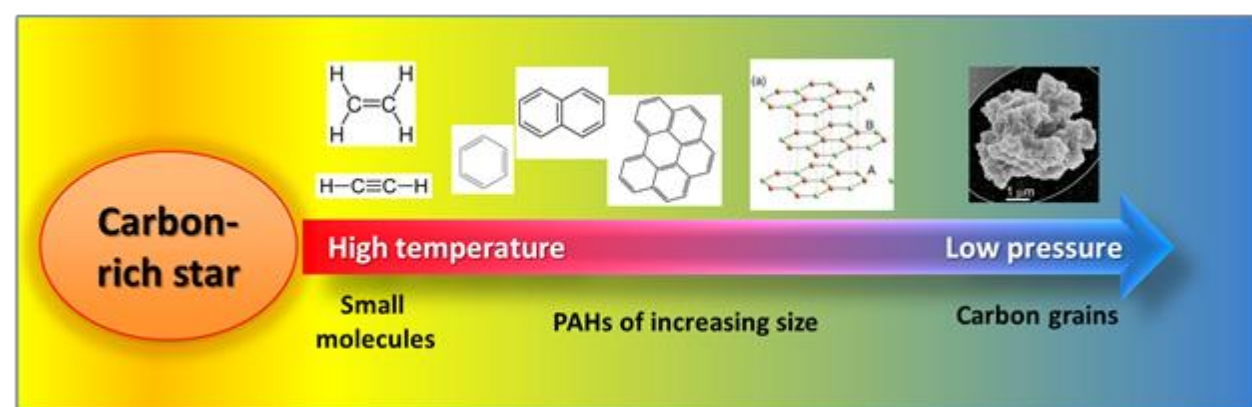


Fig. 4. Schematic representation of the formation of carbon grains in late-type carbon-rich stars.

The hydrogen-abstraction/acetylene addition (HACA) mechanism [180], involving the atomic hydrogen abstraction from benzene to phenyl radical ($\text{C}_6\text{H}_5\cdot$) [181] followed by sequential addition of two acetylene molecules, ring-closure and aromatization [182] (Fig. 5, pathway A) has been proposed to explain the large presence of naphthalene in astrochemical environments. Recently, in a combined crossed beam and theoretical study it has been demonstrated that naphthalene can be formed in the gas phase via a barrierless and exoergic reaction between a phenyl radical and vinylacetylene. This would be a low temperature alternative to the mechanism in path A, and challenges the traditional view that

PAH-formation only occurs at high temperatures [183]. This mechanism has been suggested to account for the formation of naphthalene and other PAHs in Titan's atmosphere (Fig. 5, panel B) [184].

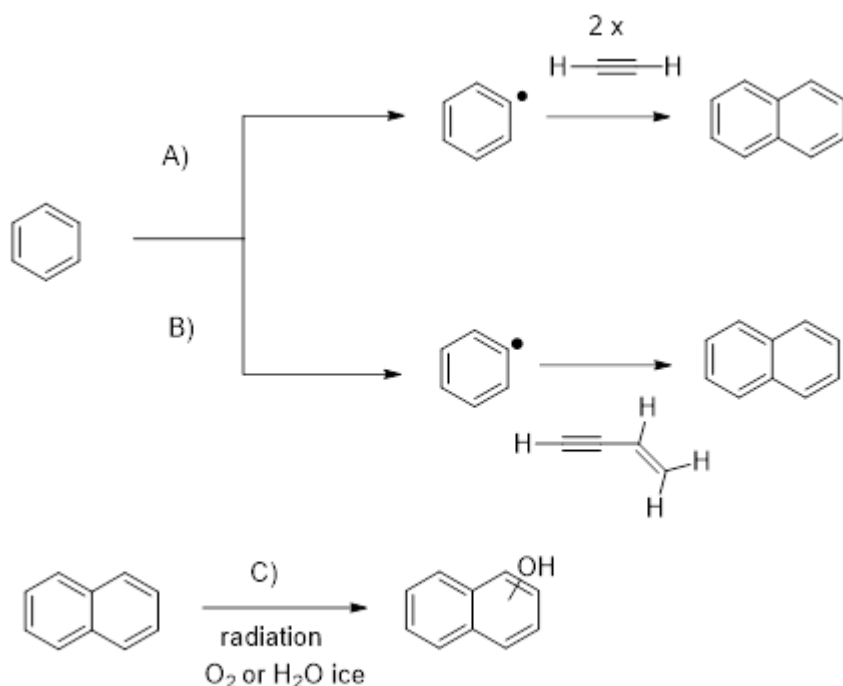


Fig. 5. Schematic representation for the radical mediated formation of naphthalene and naphthol derivatives in astrochemical environments. Pathway A: high energy addition of two acetylene molecules on a phenyl radical ($\text{C}_6\text{H}_5^\bullet$) followed by ring-closure and aromatization. Pathway B: low energy alternative mechanism based on the addition of vinylacetylene. Pathway C: general oxidation pathway for the post-synthetic oxidative transformation of the naphthalene ring mediated by the formation of HO^\bullet radicals.

Parker et al. (2012) [183] showed that the synthesis of naphthalene via this mechanism should also occur in the interstellar medium at low temperatures, including cold molecular clouds where these barrierless phenyl-type radical reactions could even lead to more complex structures like phenanthrene and anthracene [183].

The products of oxidation of PAHs, such as aromatic benzoquinone derivatives, have been initially identified in the Murchison meteorite [12]. The presence of oxidized PAHs (oxy-PAHs) in Murchison,

and presumably in other carbonaceous chondrites, is most probably due to post-synthetic processes (Fig. 5, pathway C) [177]. Both UV irradiation and cosmic rays can induce transformation of PAHs embedded in molecular cloud ice [185,186] into a wide range of reaction products [187], including oxy-PAHs derivatives and benzoquinones [188-190]; however, UV irradiation only occurs near the surfaces of molecular clouds and does not penetrate very deeply into grains, so cosmic rays are likely to be the main source of radiation and they penetrate through the whole grains inducing probably a similar chemistry to UV in ices. These molecular cloud ices are an important potential source of the organic molecules found in primitive Solar System bodies such as meteorites, asteroids, giant planet satellites, and comets [191-193]. In particular, the oxidation of naphthalene to naphthols has been reported under a large variety of experimental conditions such as UV irradiation [194] and radiolysis [195,196]. In most of these processes, reactive OH radicals play important roles, performing hydrogen abstraction processes and radical coupling reactions.

Alternatively, naphthols can be produced by thermal decomposition of simple organic salt precursors of meteorite components, as in the case of the thermal decomposition of iron oxalate. In this latter case, FTT synthesis and related processes involving CO₂, CO, H₂ and H₂O may be responsible for the observed variety of the reaction products [196], although it is relevant to note that the CI-CM-CR chondrites are generally thought to experience temperatures of no more than 150 °C. A similar reactivity has been reported in the case of thermal and photochemical processes catalyzed by clays [197], in which case naphthols can be produced from CO₂ and methane by two alternative reaction pathways: i) FTT (involving methanol) and radical ring-closure to cyclohexane followed by dehydrogenation (Fig. 6, panel A); or ii) 2+2 thermal or photoinduced cyclization of ethene molecules (Fig. 6, panel B) [198]. The actual importance of photochemical processes relevant for interstellar and early solar system processes needs however to be assessed, since photochemistry is only important at the very surface of an asteroid or comet, for instance, and only a tiny fraction of their masses can be processed in this way. Overall, these lines of evidence would indicate that the aromatic precursors of melanin, including

naphthalenes that were available in the various presolar astrochemical environments probably acted as reagents for polymerization processes that could ultimately have contributed to the formation of IOM.

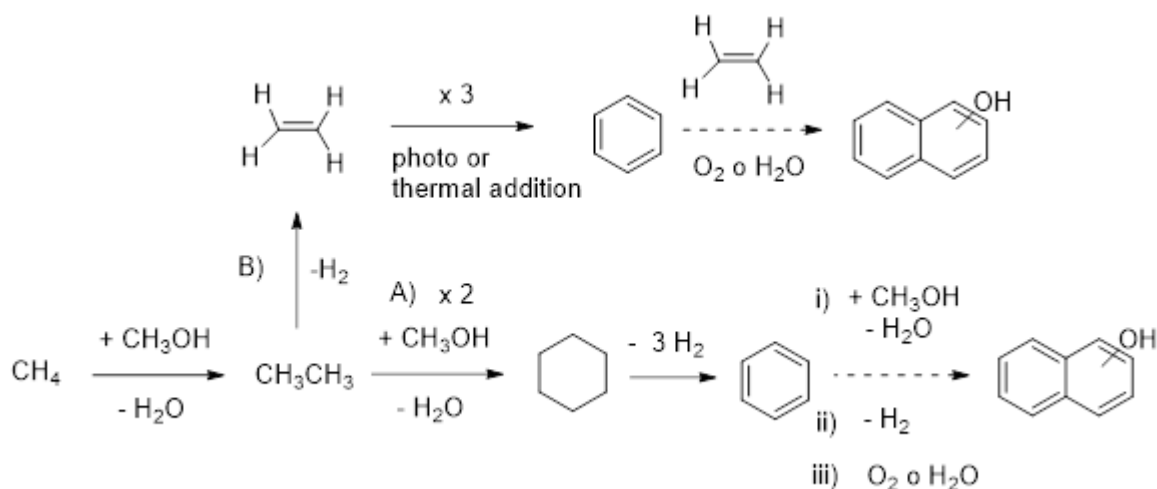


Fig. 6. Formation pathways of oxy-PHAs. A: Sequential FTT process involving methane and methanol followed by radical ring-closure dehydrogenation and terminal oxidation. B: FTT process involving methane and methanol to yield ethene followed by sequential thermal or UV cyclization and final oxidation.

In support of this view, TEM observations of carbonaceous chondrite acid residues indicated at least two morphologically distinct components of the IOM [199,200] - the major component has a fluffy texture, while the minor component occurs as solid or hollow globules/nanoglobules (also sometimes referred to as nanospheres) (Fig. 7). The fluffy material in Orgueil IOM is largely amorphous [199] but contains some regions with fringe spacings of 0.21 nm and 0.34–0.39 nm, roughly corresponding to the (100) and (002) d-spacings of graphite, respectively. Derenne et al. (2005) [201] interpreted the fringes as being local regions of order due to stacking of 2–3 of layers PAHs that are typically 2–3 rings across (i.e., ~4–9 rings/layer) [201]. However, others suggest that the fringes are the result of radiation damage in the TEM [6,143].

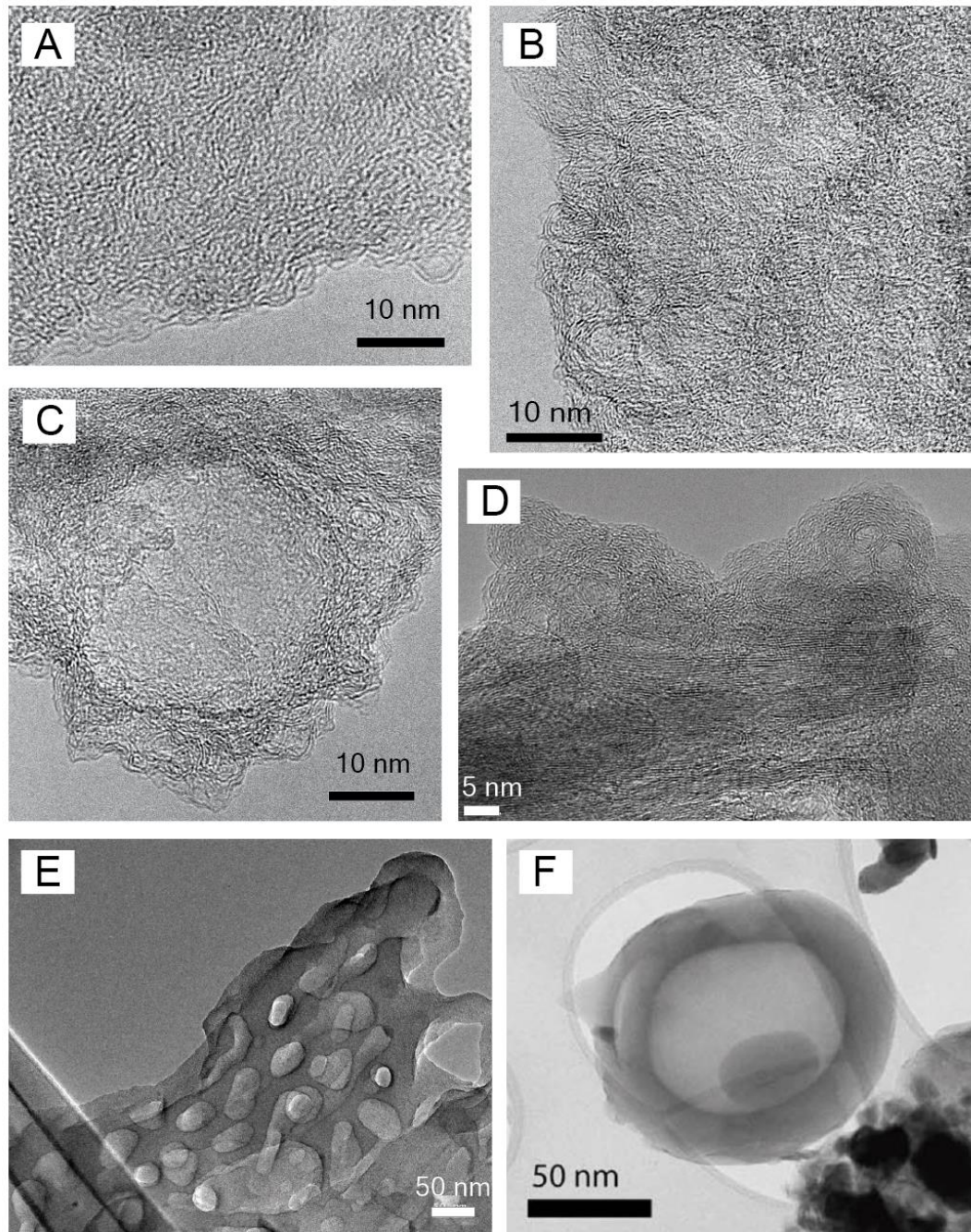


Fig. 7. TEM images revealing the various textures exhibited by IOM in chondrites A: Murchison CM2, B and C: Kainsaz CO3, D and E: Abee EH4, F: Tagish Lake. IOM usually displays a disorganized structure but can present significant porosity that can form holes and vacuoles (see C, D and E). In several type 1 and 2 carbonaceous chondrites, IOM can exhibit nanoglobules (F) which origin remains debated. Images from Le Guillou et al. (2012) [200] (A, B and C), Remusat et al. (2012) [165] (D and E) and Garvie and Buseck (2004) [134] (F).

Besides the forementioned processes in the nebula or molecular clouds, insights integrating and/or supporting current views about the role of PAHs in prebiotic processes can be gained by looking at denser environments like Titan's atmosphere, which is considered a good proxy of our primitive Earth's atmosphere. Titan, the largest moon of Saturn, has long been a subject of interest because it provides an excellent example of abiotic processing of organic material. Titan's atmosphere contains nitrogen, as major constituent, and unsaturated hydrocarbons, like acetylene, ethylene and butadiene. These are key species for the formation of PAHs and nitrogen-containing PAHs (NPAHs), the latter being important prebiotic species. In fact, they can be considered the missing link between N-bearing acyclic molecules and prebiological systems, like nucleobases. Recently, Parker and Kaiser reviewed different reaction pathways leading to the gas-phase formation of NPAHs via radical-mediated aromatization reactions [202]. For example, pyridine can be formed from the reaction of vinyl cyanide and its radical or from cyano radicals reacting with 1,3-butadiene. Furthermore, Parker and Kaiser experimentally demonstrated the feasibility and reliability of aromatization reactions occurring via neutral-neutral interactions, thus explaining how chemical evolution can proceed from acyclic systems to polycyclic aromatic hydrocarbons and showing that these are barrierless reactions that are thus feasible at low temperatures [202]. The Ion Neutral Mass Spectrometer (INMS) and the Cassini Plasma Spectrometer (CAPS) instruments onboard the Cassini spacecraft discovered the presence of a large variety of carbocations and carbanions, ranging in size from small species to heavy positive (up to 350 amu) and negative (up to 10,000 amu) ions [203]. In Ali et al. (2015) [203], the composition detected by the Cassini mass spectrometers was investigated in terms of molecular structures and the formation of complex macromolecules explained in terms of reaction mechanisms from simple molecules to complex polycyclic hydrocarbons, thus demonstrating how PAHs and NPAHs grow in size not only via neutral-neutral reactions, but also via ion-neutral reactions. Subsequently, the Nobel laureate G.A. Olah [204] extended Titan's carbocation chemistry of Ali et al. [203] to the interstellar one, thus showing a clear link between Titan's atmosphere and the interstellar medium and thus between their chemistries.

In summary, there is a lot of debate about the origin of PAHs and aromatic compounds in general. C-stars are certainly a/the major source of carbonaceous dust being added to the ISM but little is known as to how much of the dust in the ISM formed around C-stars. The current belief is that dust lifetimes in the ISM are short compared to residence times, and most dust/PAHs must form in the ISM/MCs, but how this happens is unclear.

4. Melanins as evolutionarily preserved complex multifunctional organic polymers

4.1. Classification, occurrence and physicochemical properties

Production of melanin is a universal and at the same time enigmatic form of adaptation of living organisms to the variable conditions of the Earth. The presence of various kinds of melanins in representatives of almost every large taxon suggests the evolutionary importance of melanogenesis. The traditional classification of melanins divides them into the nitrogen-containing eumelanins, pheomelanins and neuromelanins in animals, all derived from the amino acid tyrosine, and the non-nitrogenous allomelanins and pyomelanin found in plants, fungi and lower organisms that are derived mainly from a range of hydroxylated aromatic precursors [205].

From an evolutionary perspective, the appearance of nitrogen-containing units in melanin pigments would thus seem to mark the passage from unicellular or lower organisms, including fungi and plants, to the animal kingdom. At a finer level of analysis, the presence of a mutated phenotype with reddish sulfur-containing cysteine-derived melanin pigments, the pheomelanins, seems to be associated with the higher levels of evolution, including birds, mammals and humans (the red-haired Celtic phenotype). For the purposes of this paper, the main focus will be on the properties of those melanins found at the lowest evolutionary levels that may be closer to the prebiotic precursors, especially those in bacteria and fungi. For exhaustive coverage of the topic the reader is referred to several recent reviews [206-209].

Synthesis of melanins in bacteria was initially thought to occur only through the DOPA pathway. In the seventies a brown pigment derived from homogentisic acid was identified in *Pseudomonas*

aeruginosa and was referred to as pyomelanin (Fig. 8) [210]. The same type of pigment was later identified in *Streptomyces avermitilis* [211] and again in *Pseudomonas putida* [212,213].

Homogentisate-derived melanin formation was described in mutated *Vibrio cholerae* strains [214] and in mutated *Burkholderia cepacia* complex isolates [215].

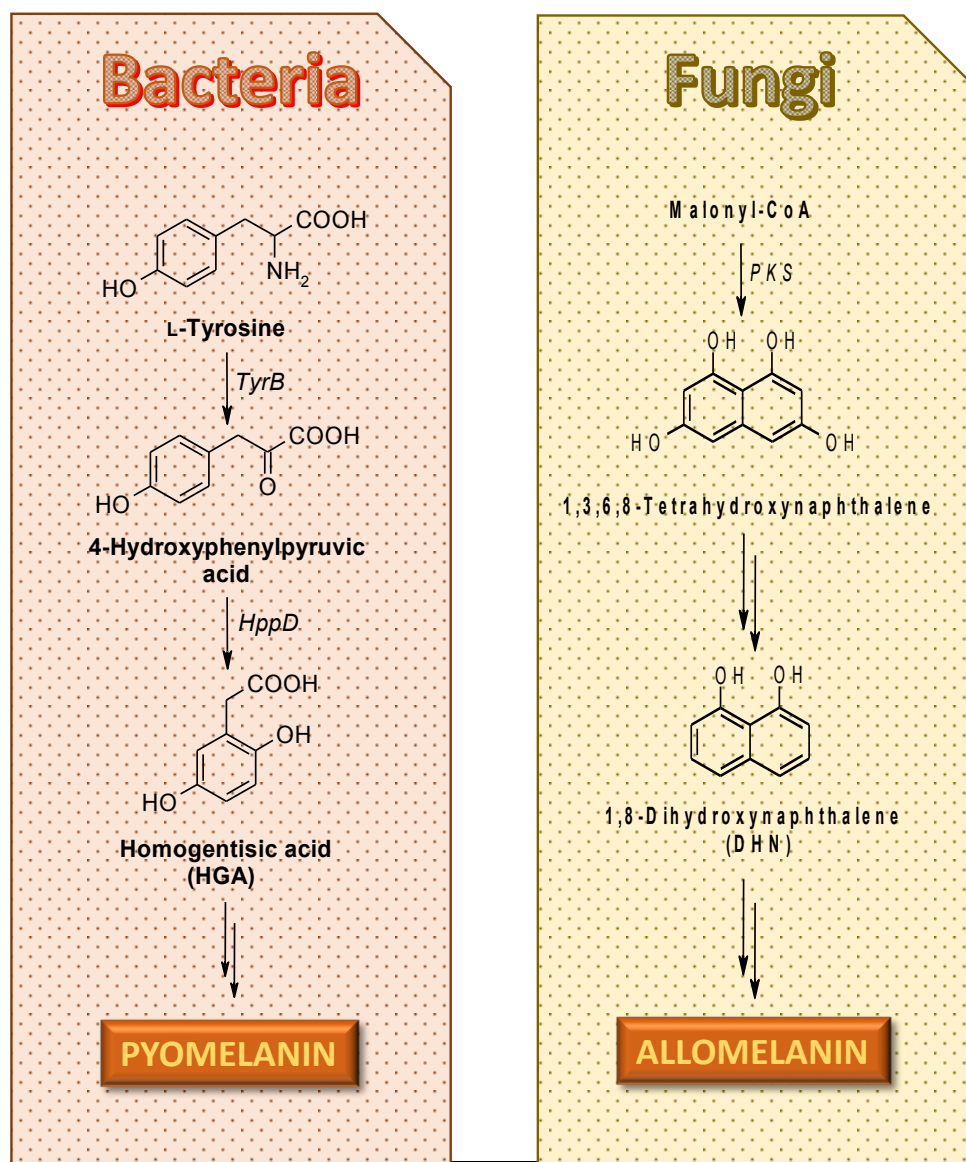


Fig. 8. Schematic view of the main melanin biosynthetic pathways in bacteria and fungi. Minor pathways have been omitted for the sake of simplicity.

In fungi, melanins arise mainly via two pathways (Fig. 8) [208]. The main one stems from the oxidative polymerization of 1,8-dihydroxynaphthalene (DHN), a metabolite produced via a polyketide synthase (PKS) pathway [216-219]. The polymerization process may be brought about by multicopper

oxidases in *Asperigillus fumigatus* [220,221], *Cochliobolus heterostrophus* [222] and *Cochliobolus orbiculare* [223]. A few fungi synthesize melanins starting from tyrosine, via the oxidative polymerization of its metabolites DOPA and homogentisic acid, this latter leading to the formation of the brownish pyomelanins [224,225].

Herein, for the sake of simplicity, all non-nitrogenous melanins from bacteria and fungi will be collectively referred to as allomelanins. Elemental composition data for representative examples of bacterial and fungal melanins is reported in Table 2.

Table 2. Comparative view of elemental compositions of bacterial and fungal melanins

	C (%)	H (%)	O (%)	N (%)	S (%)	Reference
Fungal melanin						
<i>Cryptococcus neoformans</i>	58.99	7.11	28.34	5.56	-	[226]
<i>Spissiomycetes endophytica</i>	52.69	4.69	42.15	0.47	-	[227]
<i>Tuber melanosporum</i>	47.2	6.0	45.2	1.0	< 0.3	[228]
<i>Auricularia auricula</i>	41.18	5.56	51.60	1.66	-	[226]
<i>Termitomyces aluminosus</i>	54.68	3.54	26.92	2.49	12.36	[229]
<i>Tuber melanosporum</i> Vitt.	52.10	4.20	28.31	6.16	9.23	[230]
<i>Boletus griseus</i>	56.38	5.86	28.04	6.17	2.44	[231]
Bacterial melanin						
<i>Klebsiella sp. GSK</i>	47.9	6.9	-	12.0	0.9	[232]
<i>Bacillus subtilis</i>	70.0	-	-	12.0	4.1	[233]

<i>Bacillus subtilis</i>	41.7	6.56	-	8.57	0.98	[234]
<i>4NP-BL</i>						

Studies over the last decades have clarified some important chemical and physical properties of allomelanins [207] that are common to all members of this class of polymers regardless of structural and biosynthetic differences [235]. These properties include, among others: *a*) a black or dark coloration due to broadband, often featureless, absorption spectra covering the entire UV-visible range [236]; *b*) a marked structural heterogeneity; *c*) almost complete insolubility in water and organic solvents; *d*) an intrinsic paramagnetic character responsible for a broad stable EPR signal [237]; *e*) semiconductor-type electrical conductivity properties [238]; *f*) efficient non-radiative excited state deactivation; *g*) a redox behavior associated with antioxidant activity; *h*) a strong metal cation binding and drug affinity, and *i*) efficient gas adsorption properties.

Most of these properties, especially *a*) and *d*-*g*), are evidently associated with the unusual chemical complexity of all melanin-type polymers at the level of the conjugated π -system. Such a complexity stems from various interrelated levels of structural disorder, including monomer disorder, molecular weight disorder, isomerism-related disorder, redox disorder and supramolecular disorder [236,237].

4.2. Biological roles and functions

The ability to produce melanin from various phenolic substrates is typical of various microorganisms; however, although the subject of multidisciplinary investigations since time immemorial, the biological roles of melanins are still a matter of debate. Besides photoprotection, the accepted functions of most melanins include camouflage, antioxidant and free radical scavenger defense, UV dissipation, ionic conductor behavior, and de-toxification via binding of metals and xenobiotics [239]. These diverse functions suggest that the utility of melanogenesis is rooted in the peculiar physical and chemical characteristics of the pigments. Indications of the past biological role of melanin pigments can be obtained from the occurrence of allomelanins in several bacteria and fungi that appeared on Earth before plants and higher organisms. Black melanin pigments in bacteria play diverse biological

roles, including protection from UV radiation, heavy metals and oxidative stress, and have been implicated in pathogenesis and survival in adverse environments (Fig. 9) [206].

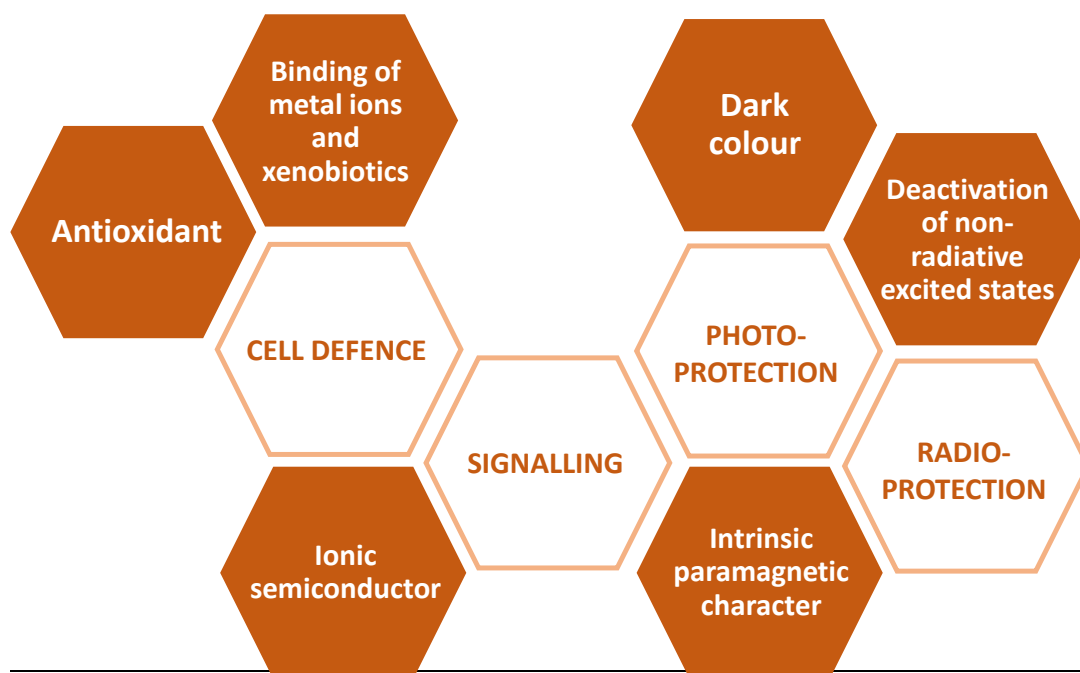


Fig.9. The conceptual interconnections between main properties (full hexagons) and biological functions (empty hexagons) of melanins.

In *Bacillus anthracis*, pigmented mutants are more resistant to UV exposure [240], and melanized spores of *Streptomyces griseus* are more resistant to UV irradiation than colorless mutants [241]. In addition, melanized *Pseudomonas aeruginosa* strains are more tolerant to oxidative stress [242]. Interestingly, protection of bacterial melanin from photodamage can be extended to the organisms living in close association, as in the case of the darkly pigmented sponges that contain *Vibrio*, *Providencia*, and *Shewanella* symbionts [243]. The important role of melanins is supported by the fact that several non-phylogenetically related bacteria sharing the same ecological environment produce the pigment.

In fungi melanin is an important protective factor against the adverse effects of environmental stresses, such as UV radiation, heat, cold, drying, salts, heavy metals and radionuclides [244]. Melanins are believed to account for the ability of fungi to survive high levels of ionizing radiation [245,246]. Under

extreme conditions, mycobiota usually produce higher proportion of melanized fungi, such as in the case of radionuclide contamination. Likewise, the amount of melanized spores is higher in air exposed environments than in soils and dark-colored spores are more resistant to UV irradiation. Microcolonial yeast-like fungi that grow on rocks under the conditions of extreme temperatures, e.g. in the Antarctic, or drought and low concentrations of organic compounds, are strongly melanized. On this basis, it can be concluded that melanin production in fungi represents an example for evolutionary convergence, being produced at the end of different metabolic pathways, even in one taxon, as a means of achieving efficient adaptation. The possible evolutionary origins of melanogenesis have been suggested to include extracellular oxidative polymerization of phenolic compounds caused by enzymes supporting this process, followed by a gradual adaptation to the intracellular control of the process.

The question of melanin's radiation protective function can be also looked at from a paleobiological perspective [209]. Highly melanized fungal spores have been detected in deposits of the early Cretaceous period when the Earth crossed 'magnetic zero' with decreased defense against solar and stellar radiation [247]. In this scenario, the dying out of many species of animals and plants was hypothesized to involve an unseen solar companion star, a hypothetical red dwarf or brown dwarf generally referred to as Nemesis, triggering periodic mass extinction, although this hypothesis has been questioned [248]. Melanin-related virulence may have also been implicated in the possible proliferation of fungal species that contributed to the extinctions at the end of the Cretaceous period [249]. In concluding this section, we concur with Dadachova et al., 2007 [209] that: "*Melanins are ancient pigments that are likely to have been selected in evolution because they enhanced the survival of microbes in both the environment and in various hosts. The emergence of melanin as a nonspecific radioprotective material may be a result of the relative ease with which these complicated aromatic structures can be synthesized from a great variety of precursors.*"

5. The role of IOM in prebiotic processes: from the melanin connection to the PriME scenario.

5.1. Structural and physicochemical similarities between IOM and allomelanin

As mentioned previously, a point-by-point comparison between the IOM and allomelanin should provide the reader with the proper conceptual basis to support the PriME scenario. However, before addressing this issue it seems necessary to carefully consider the following caveats:

- 1) it may sound incongruous to compare petrified organic mixtures from different chondritic samples, that experienced a variety of astrophysical conditions, with the products of genetically- and enzymatically controlled biosynthetic pathways like the allomelanins.
- 2) a direct comparison of elemental compositions and other structural data is likewise difficult within the melanin group, considering: a) the presence in natural pigments of the biological matrix possibly affecting most analytical data, b) the broad variety of purification protocols for natural pigments, making reproducibility an arduous task (use of synthetic melanins is generally preferred for comparative purposes), and c) the wide range of melanin pigments with different compositions, which would make an *a priori* choice of the reference group both questionable and arbitrary.

Bearing in mind these caveats, it is noteworthy from the preceding overview that *chondritic IOM shares fundamental physico-chemical properties with allomelanin-type materials*. More generally, it appears that structurally disordered organic polymers from phenolic precursors share a number of properties in common that partially override their intrinsic (bio)synthetically-determined structural diversity. These properties, or more properly, *emergent properties*, reflect strong intermolecular interactions between the highly heterogeneous molecular components that determine basic aggregation processes underlying the amorphous morphologies of most of these polymers and eventually account for their complex natures and exotic features. Evidence in support of the analogy between IOM and allomelanins comes mainly from elemental composition data, solid state ^{13}C NMR, EPR and X-ray spectroscopy, and TEM analysis.

A comparison of the EPR signal features of IOM from Murchison, Orgueil, Paris, Tagish lake (g-factor values centered around 2.0031) and of allomelanins from *Aspergillus niger* and *Aspergillus fumigatus* (g-factor values of 2.0035) would suggest similar carbon-centered radicals. This observation

corroborates the view that strong aggregation and intermolecular perturbation of π -electron systems in the solid state generates chemical complexity that can produce emergent properties from self-assembly and self-organization as universal features of primitive disordered organic polymers. An overview of the unifying set of properties and characteristics of IOM and allomelanins is given in Fig. 10.

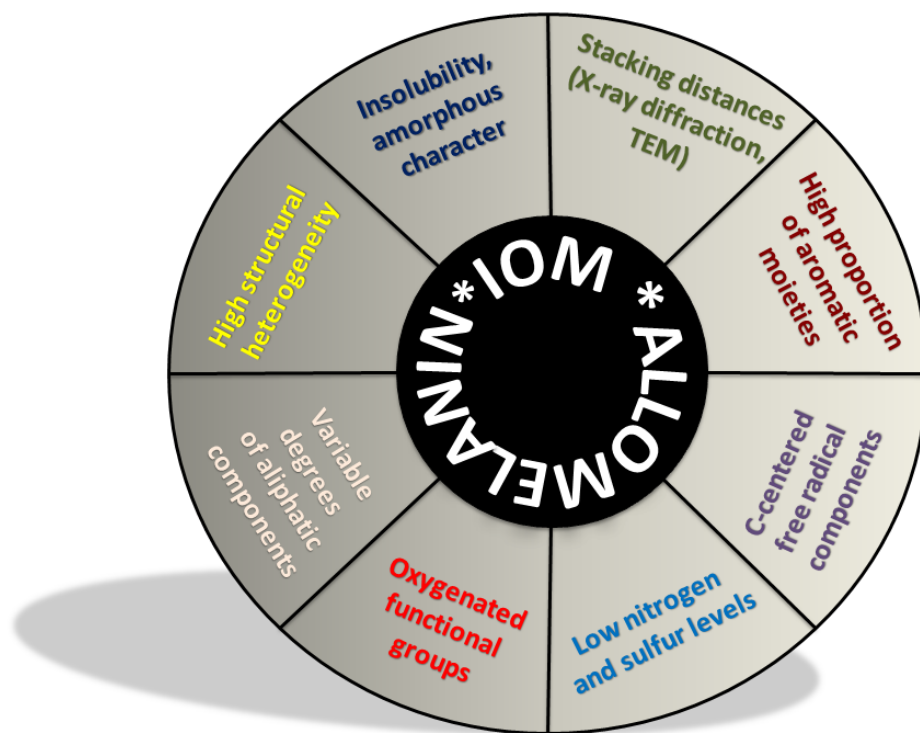


Fig. 10. Unifying set of properties and characteristics of IOM and allomelanins.

The above similarities, coupled with the known properties and functions of allomelanins and most other complex phenolic polymers, provide the basic chemical and physical ground for inferring the possible role of IOM in abiogenesis and the origin of important prebiotic molecules.

5.2. Chemical requirements at the origin of life

The origin of life and of life-related molecules, be it on the Earth or other planetary or non-planetary contexts, has been suggested to depend on at least eight critical reaction conditions: (1) reductive gas phase, (2) alkaline pH, (3) freezing temperature, (4) water, (5) dry/dry-wet cycle, (6) coupling with

high energy reactions, (7) heating-cooling cycle in water, and (8) input of life building blocks from, e.g., carbonaceous chondrites [250], comets [102], and micrometeorites [251].

Connected with these reaction conditions are three different sets of basic requirements that are schematically illustrated in Fig. 11.

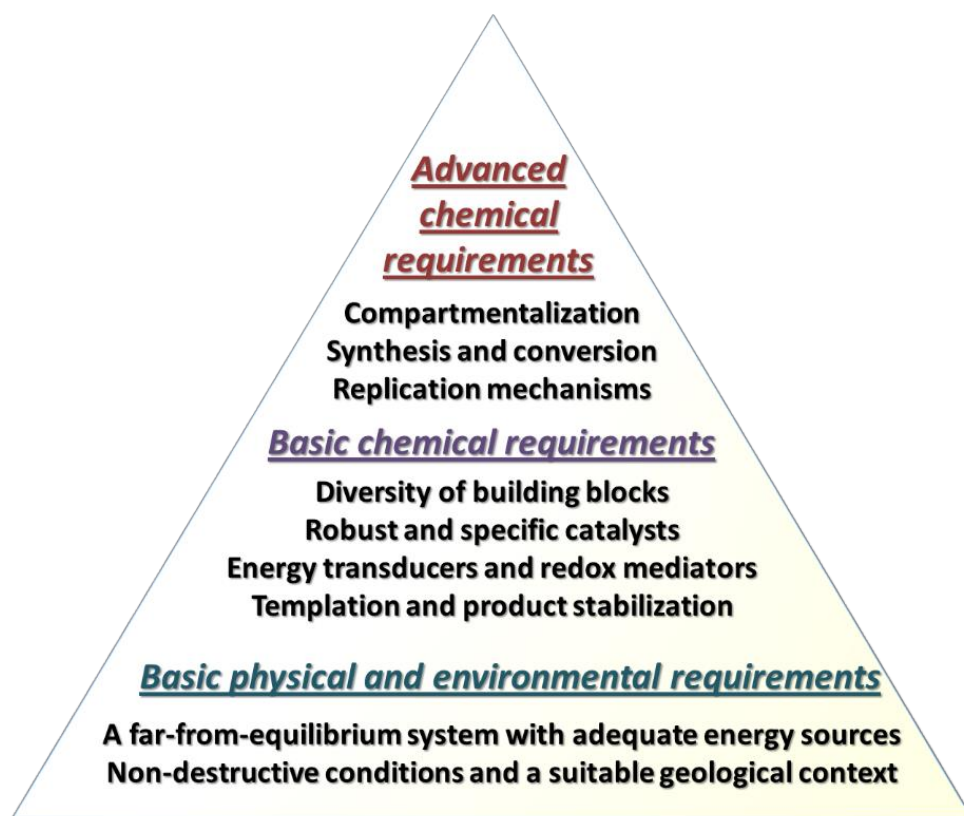


Fig. 11. Main physical and chemical requirements for the origin of life

A reductionistic and unifying scenario that could fulfill most of these essential requirements for the origin of life and the need for efficient energy dissipation must include a number of requisites such as:

- a) provision of all the elements and essential life molecules building blocks; b) a versatile catalytic ability; c) efficient energy transducing properties in a far-from-equilibrium system; d) redox behavior; e) distinct aromatic domains providing templating motifs for biomolecule, e.g., nucleobase or porphyrin, self-assembly; f) the capacity to protect life-molecules and evolving protocells from damaging radiation and other destructive agents; g) the ability to ensure efficient compartmentalization

for the assembly of protocells, minimal cells and other microenvironments supportive of life; h) facile replication mechanisms.

Based on the data reported in the present review, it is reasonable to hypothesize that IOM, by virtue of its similarities with allomelanins may have served most of the functions expected for the PriME in abiogenesis, as briefly summarized below.

- a) IOM, like all melanins and phenolic polymers, incorporates, besides carbon, various oxygenated, nitrogenous and sulfur-based groups that may have produced small, reactive molecules upon UV, chemical or thermal degradation that were useful in the development of early life [252-254].
- b) IOM, by virtue of its structural similarity with melanin-type polymers, could have provided a catalytic environment for various processes (an example of catalytic properties of melanins is provided in [255]).
- c) IOM may have played a role in energy transducing mechanisms (see e.g., melanin role in organic electrochemical transistor (OECT) experiments [256]). In this regard, IOM first and melanin later might have been crucial players in Michaelian's (2017) thermodynamic dissipation perspective [257]. In this latter, the biotic-abiotic co-evolution of the biosphere was envisaged as being based on the observation that the photon dissipation efficacy of Earth in its solar environment has increased ever since the formation of Earth in the Hadean and that this evolution has incurred the dissipative structuring of pigments: *"...Their proliferation to far beyond equilibrium concentrations can be understood in terms of non-linear, non-equilibrium thermodynamic principles directing autocatalytic photochemical reactions in which these pigments catalyze the dissipation of the same thermodynamic potential (the solar photon flux) that produced them."*
- d) IOM can serve as a protective support for biomolecule assembly, preventing their photodegradation. Indirect evidence supporting this hypothesis comes from the known quenching ability of melanins toward photoexcited states [258].

- e) IOM can be implicated in very primitive and non-classic compartmentalization and replication processes at the origin of life. For example, TEM observations [200] revealed small cavities formed by defects in aromatic plane stacking that may be implicated in compartmentalization phenomena and specific confinement of chemical transformations. A relevant observation suggesting possible control over the interplay of metabolic and replication processes relates to the ability of non-polar molecules extracted from the Murchison carbonaceous chondrite to form vesicles structure. Reactions in liquid crystalline organizing matrices have also been proposed as a promising approach for the way primitive cells could have emerged on the early Earth [259].

5.3. The PriME scenario and the Universal Multifunctional Adaptive Mediator (UMAM) paradigm.

The preceding sections identified three main pillars for the “Primordial Multifunctional Organic Entities -PriMEs” scenario: a) IOM can be regarded as an archetypal abiotic melanin sharing key properties in common with fungal melanins and other forms of primitive organic matter, like kerogens; b) melanins are evolved, specialized and genetically-determined biopolymers that play a broad variety of roles in an adaptive and evolutionarily-determined manner; and (by implication): c) IOM from chondrites and other sources may provide a missing universal multifunctional mediator in the early Earth because of its melanin-like properties, whereby it can fulfill, or contribute to fulfill, most of the chemical requirements for the transition from an inorganic world to a prebiotic world producing structurally defined biomolecules.

In this scenario, IOM may well be a source of PriMEs operating at the crossroads of various prebiotic processes, because of:

- a) a hybrid robust structure deeply incorporated into the stony matrix following extensive thermal processing;

- b) a broad range of chemically diverse functional groups and life-relevant elements, from nitrogen to oxygen and sulfur;
- c) possible melanin-type physicochemical properties, including energy absorption, transduction and dissipation, a free radical character, catalytic activity, high adsorption capacity for gases and organic substances, metal chelation [260]

The striking analogy of IOM with melanin biopolymers, especially allomelanins from lower organisms, may be finally integrated into the novel paradigm of the Universal Multifunctional Adaptive Mediator (UMAM, see Fig. 12) that has been preserved and transformed through the development of the biosynthetic machinery for melanins to fulfill the manifold requirements for abiogenesis. From this perspective, melanins can be considered as a genetically/biochemically controlled evolution of the IOM/PriME.

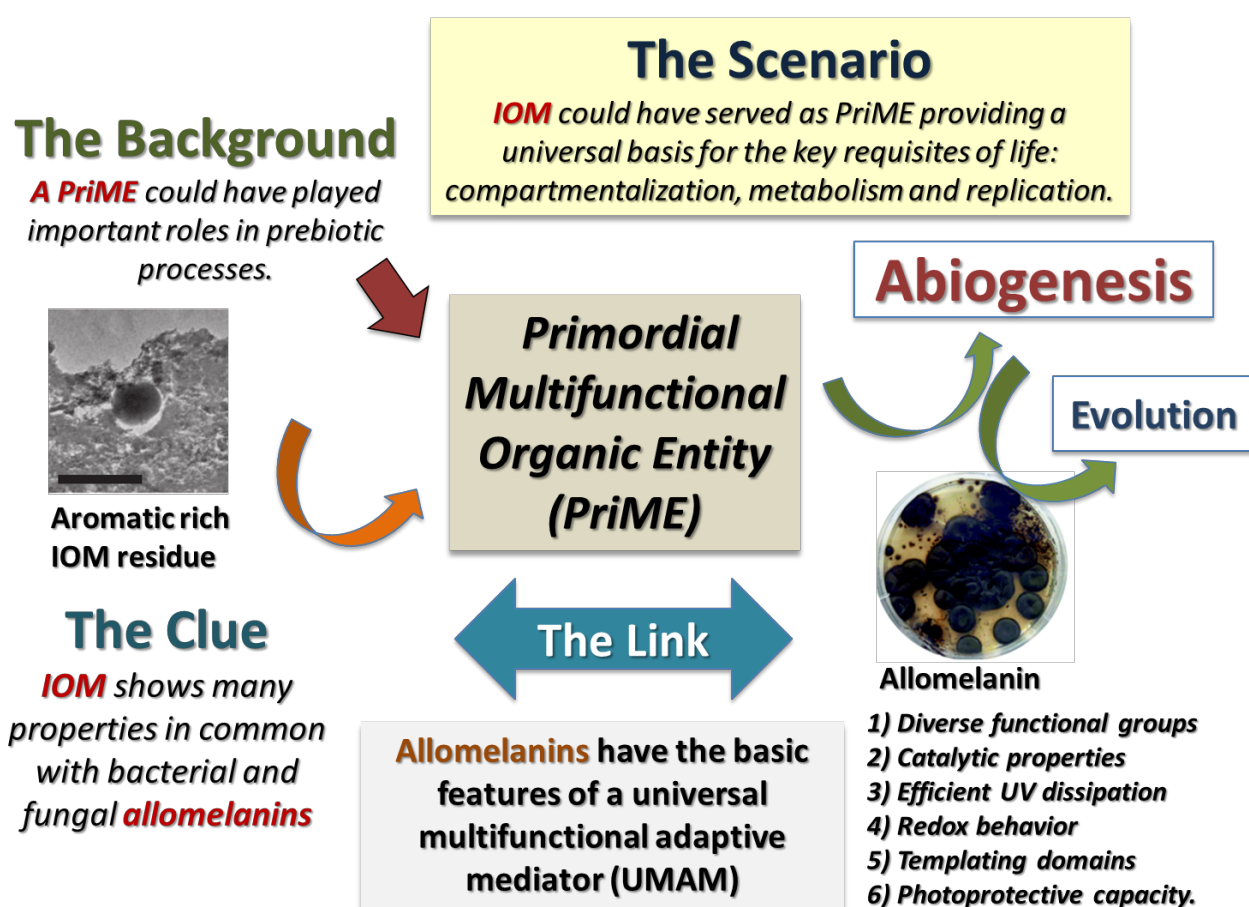


Fig. 12. The conceptual background of the PriME scenario and the UMAM paradigm.

Some of the most stimulating challenges that should be addressed in ongoing and future studies include: a) defining in more detail the structure and origin of IOM, through the use of synthetic models, b) probing the potential role of IOM as prebiotic mediator via a solid experimental approach, and integrating the emerging findings about IOM into the most advanced models and theories about the origin of life. Important issues in this connection include:

- 1) the mechanisms of conversion of IOM into small building blocks that would be useful in the development of life, a process that is probably related to the exposure of chondrites to energetic radiation, discharges, thermal shocks and/or aqueous environments at the surface of young planets;
- 2) the demonstration that the organisms closer to the roots of the tree of life use melanins, for instance, the organisms that were responsible for 3-3.5 billion years old stromatolites or for the rise of O at 2.5 billion years.

5.4. The PriME scenario and Virtual Reality

A promising approach to inquire into the role of PAHs in the formation of IOM and as both a template and a mediator in key prebiotic processes would be based on advanced computational tools like artificial intelligence (AI) and immersive virtual reality (IVR) coupled to powerful hardware and software for numerical simulations. A new perspective is offered by the transition from the macroscopic to the microscopic world within a virtual reality framework thanks to the extension of a very general immersive IVR tool (Unity-3D [261]) to the molecular world (Unity-Mol [262]) and the development of commodity hardware with remarkable performance (e.g., Vive [263] or Oculus Rift [264]). In particular, a general tool, named Tardis, was developed, which can be employed for the exploration of chemical events occurring in the interstellar space, on exoplanets or meteorites and on the primordial Earth [265-267]. 3D representations of putative aromatic building blocks of IOM and of fungal melanin (Fig. 13) may be useful in this regard as a basis to enquire into the physical and chemical properties of these materials once properly framed into a IVR prebiotic scenario.

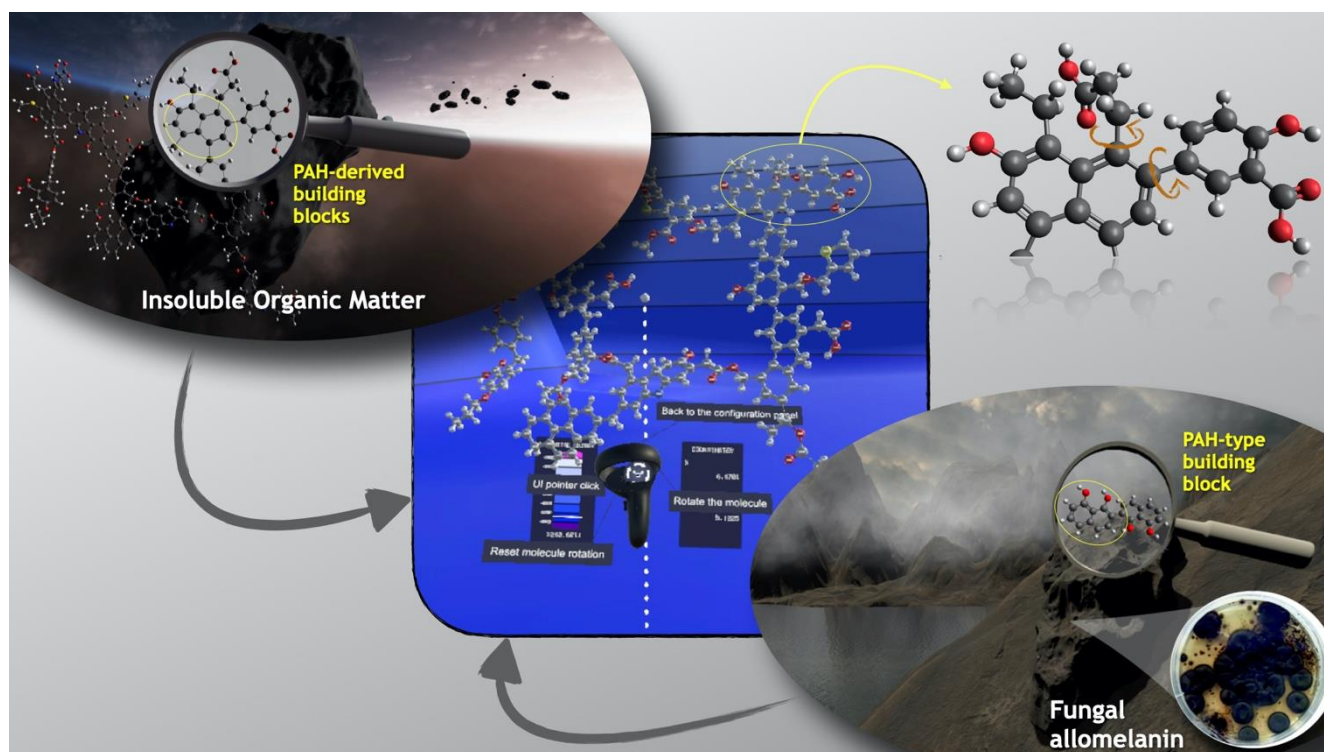


Fig. 13. Transition from the macroscopic to the molecular world in two representative scenarios: PAH-derived building blocks on a carbonaceous meteorite and on fungi on the Earth. After the molecular system is perceived, two torsional degrees of freedom have been selected and the corresponding PES generated and analyzed in an IVR context.

Virtual reality may indeed allow inspection of a specific asteroid by opening the metadata associated to it, then shifting to the microscopic level, where the structure and properties of the minerals composing the asteroid and, possibly, a chemisorbed molecule like a PAH, can be investigated. As an alternative scenario, reactions involving PAHs can be visualized and investigated in the virtual environment of an exoplanet, where hypothetical fungi could evolve and produce, *inter alia*, allomelanins. Then, starting from an automatic perception of the molecular system, it is possible to analyze structural and conformational properties thanks to the effective integration of AI and IVR tools. For instance, by properly reducing the dimensionality of the potential energy surface (PES) to two coordinates and building a three-dimensional representation where the selected coordinates are used as the latitude and the longitude and the associated value of the PES as the elevation, it is possible

to explore the conformational behavior (see Fig. 13) or the reactivity of the target system or of its fragments by an immersive exploration, in which the user can walk around a virtual world having the PES as terrain.

It is expected that new exciting opportunities can derive from the latest developments of immersive virtual reality approaches as versatile tools for aiding multidisciplinary research like those at the origin of life.

6. Conclusions and perspectives

The present survey developed from the realization that the IOM from chondrites and the allomelanins found in lower organisms display similar physicochemical properties. Considering that: *a*) spectral evidence suggests the presence of PAH-related aromatic domains in chondritic IOM; *b*) PAHs and related aromatic species are believed to be some of the primary components of complex organic molecules in the interstellar medium; *c*) PAH-type systems, like DHN, are the key building blocks of allomelanins, and *d*) allomelanins, like DHN melanin, exhibit physical features and chemical properties that may be compatible with a role in primitive compartmentalization, catalysis and replication phenomena, it can be argued that IOM delivered in favorable environments by chondrites may have played a role as a PriME, promoting key prebiotic processes by means of its PAH-derived composition.

The PriME scenario and other commonly accepted prebiotic scenarios are not mutually exclusive. It is therefore important to test prebiotic models based on the PriME scenario by addressing a number of pending issues, including in particular the following:

- 1) the occurrence, origin and role of PAHs in IOM from chondrites;
- 2) the ability of IOM from chondrites to behave as multifunctional allomelanin-like systems playing key roles in putative prebiotic processes;

- 3) the processes underlying the postulated transition from chemically-controlled PAH polymerization in IOM and related materials delivered to the early Earth to the genetically and enzymatically-controlled allomelanin synthesis from DHN in early living organisms.

Pursuit of these issues at interdisciplinary level is expected to provide significant advancements into the nature of prebiotic processes with a new focus on PAH-derived species.

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